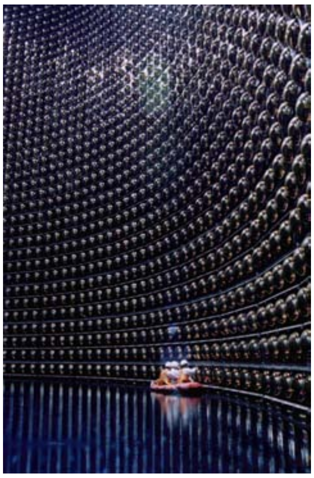
Particle Physics

Michaelmas Term 2011
Prof Mark Thomson



Handout 11 : Neutrino Oscillations

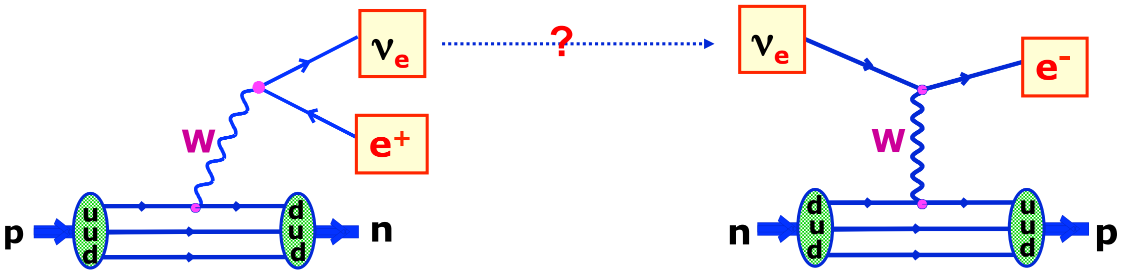
Neutrino Flavours Revisited

★ Never **directly** observe neutrinos – can only detect them by their weak interactions. Hence by **definition** ν_e is the neutrino state produced along with an electron. Similarly, charged current weak interactions of the state ν_e produce an electron

$$\boxed{\nu_e, \nu_\mu, \nu_\tau \text{ = weak eigenstates}}$$

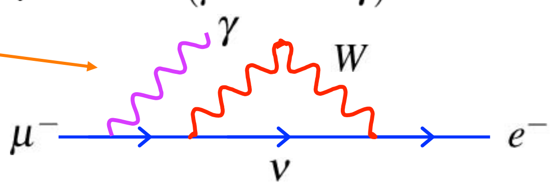
★ For many years, assumed that ν_e, ν_μ, ν_τ were massless fundamental particles

- **Experimental evidence:** neutrinos produced along with an electron always produced an electron in CC Weak interactions, etc.



- **Experimental evidence:** absence $\mu^- \to e^- \gamma$ $\quad \mathrm{BR}(\mu^- \to e^- \gamma) < 10^{-11}$

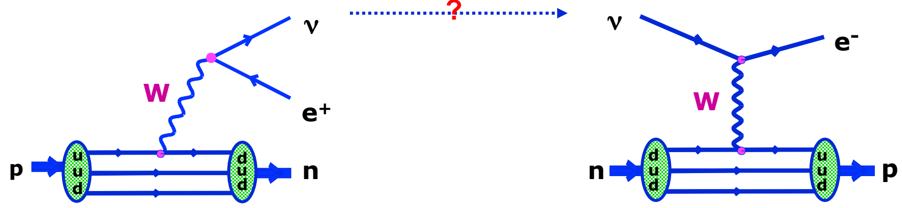
Suggests that ν_e and ν_μ are distinct particles otherwise decay could go via:



Mass Eigenstates and Weak Eigenstates

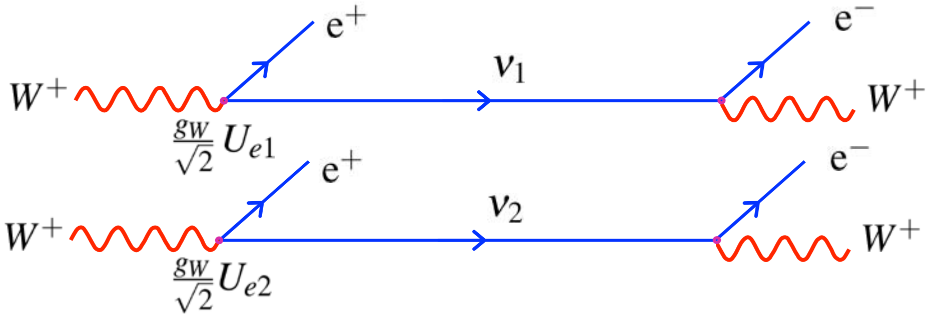
★ The essential feature in understanding the physics of neutrino oscillations is to understand what is meant by weak eigenstates and mass eigenstates ν_1, ν_2

★ Suppose the process below proceeds via two fundamental particle states



i.e.

and



★ Can't know which **mass eigenstate** (fundamental particle ν_1, ν_2) was involved

★ In Quantum mechanics treat as a coherent state $\psi = \nu_e = U_{e1}\nu_1 + U_{e2}\nu_2$

★ ν_e represents the wave-function of the coherent state produced along with an electron in the weak interaction, i.e. the weak eigenstate

Neutrino Oscillations for Two Flavours

★ Neutrinos are produced and interact as weak eigenstates, ν_e, ν_μ

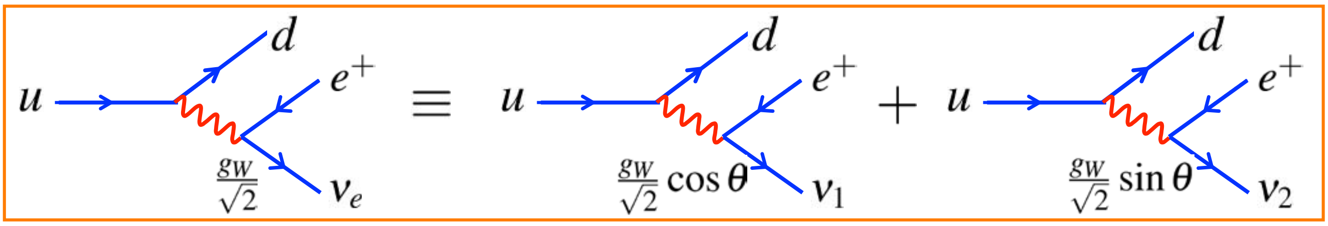
★ The weak eigenstates as coherent linear combinations of the fundamental "mass eigenstates" ν_1, ν_2

★ The mass eigenstates are the free particle solutions to the wave-equation and will be taken to propagate as plane waves

$$|\nu_1(t)\rangle = |\nu_1\rangle e^{i\vec{p}_1.\vec{x}-iE_1t} \qquad |\nu_2(t)\rangle = |\nu_2\rangle e^{i\vec{p}_2.\vec{x}-iE_2t}$$

★ The weak and mass eigenstates are related by the unitary 2x2 matrix

$$\begin{pmatrix} \nu_e \\ \nu_\mu \end{pmatrix} = \begin{pmatrix} \cos\theta & \sin\theta \\ -\sin\theta & \cos\theta \end{pmatrix} \begin{pmatrix} \nu_1 \\ \nu_2 \end{pmatrix} \qquad (1)$$



★ Equation (1) can be inverted to give

$$\begin{pmatrix} \nu_1 \\ \nu_2 \end{pmatrix} = \begin{pmatrix} \cos\theta & -\sin\theta \\ \sin\theta & \cos\theta \end{pmatrix} \begin{pmatrix} \nu_e \\ \nu_\mu \end{pmatrix} \qquad (2)$$

• Suppose at time $t=0$ a neutrino is produced in a pure ν_e state, e.g. in a decay $u \rightarrow de^+\nu_e$

$$|\psi(0)\rangle = |\nu_e\rangle = \cos\theta|\nu_1\rangle + \sin\theta|\nu_2\rangle$$

• Take the z-axis to be along the neutrino direction

• The wave-function evolves according to the time-evolution of the mass eigenstates (free particle solutions to the wave equation)

$$|\psi(t)\rangle = \cos\theta|\nu_1\rangle e^{-ip_1.x} + \sin\theta|\nu_2\rangle e^{-ip_2.x}$$

where $p_i.x = E_i t - \vec{p}_i.\vec{x} = E_i t - |\vec{p}_i|z$

• Suppose the neutrino interacts in a detector at a distance **L** and at a time **T**

$$\phi_i = p_i.x = E_i T - |\vec{p}_i|L$$

gives $$\boxed{|\psi(L,T)\rangle = \cos\theta|\nu_1\rangle e^{-i\phi_1} + \sin\theta|\nu_2\rangle e^{-i\phi_2}}$$

★ Expressing the mass eigenstates, $|\nu_1\rangle, |\nu_2\rangle$, in terms of weak eigenstates (eq 2)

$$|\psi(L,T)\rangle = \cos\theta(\cos\theta|\nu_e\rangle - \sin\theta|\nu_\mu\rangle)e^{-i\phi_1} + \sin\theta(\sin\theta|\nu_e\rangle + \cos\theta|\nu_\mu\rangle)e^{-i\phi_2}$$

$$\boxed{|\psi(L,T)\rangle = |\nu_e\rangle(\cos^2\theta e^{-i\phi_1} + \sin^2\theta e^{-i\phi_2}) + |\nu_\mu\rangle\sin\theta\cos\theta(-e^{-i\phi_1} + e^{-i\phi_2})}$$

★ If the masses of $|\nu_1\rangle, |\nu_2\rangle$ <u>are the same</u>, the mass eigenstates remain in phase, $\phi_1 = \phi_2$, and the state remains the linear combination corresponding to $|\nu_e\rangle$ and in a weak interaction will produce an electron

★ If <u>the masses are different</u>, the wave-function no longer remains a pure $|\nu_e\rangle$

$$\begin{aligned} P(\nu_e \rightarrow \nu_\mu) &= |\langle\nu_\mu|\psi(L,T)\rangle|^2 \\ &= \cos^2\theta\sin^2\theta(-e^{-i\phi_1} + e^{-i\phi_2})(-e^{+i\phi_1} + e^{+i\phi_2}) \\ &= \frac{1}{4}\sin^2 2\theta(2 - 2\cos(\phi_1 - \phi_2)) \\ &= \sin^2 2\theta \sin^2\left(\frac{\phi_1 - \phi_2}{2}\right) \end{aligned}$$

★ The treatment of the phase difference

$$\Delta\phi_{12} = \phi_1 - \phi_2 = (E_1 - E_2)T - (|p_1| - |p_2|)L$$

in most text books is dubious. Here we will be more careful…

★ One could assume $|p_1| = |p_2| = \mathrm{p}$ in which case

$$\Delta\phi_{12} = (E_1 - E_2)T = \left[(\mathrm{p}^2 + m_1^2)^{1/2} - (\mathrm{p}^2 + m_2^2)^{1/2}\right]L \qquad L \approx (c)T$$

$$\Delta\phi_{12} = \mathrm{p}\left[\left(1+\frac{m_1^2}{\mathrm{p}^2}\right)^{1/2} - \left(1+\frac{m_2^2}{\mathrm{p}^2}\right)^{1/2}\right]L \approx \frac{m_1^2 - m_2^2}{2\mathrm{p}}L$$

★ **However we have neglected that fact that for the same momentum, different mass eigenstates will propagate at different velocities and be observed at different times**

★ **The full derivation requires a wave-packet treatment and gives the same result**

★ **Nevertheless it is worth noting that the phase difference can be written**

$$\Delta\phi_{12} = (E_1 - E_2)T - \left(\frac{|p_1|^2 - |p_2|^2}{|p_1| + |p_2|}\right)L$$

$$\Delta\phi_{12} = (E_1 - E_2)\left[T - \left(\frac{E_1 + E_2}{|p_1| + |p_2|}\right)L\right] + \left(\frac{m_1^2 - m_2^2}{|p_1| + |p_2|}\right)L$$

★ **The first term on the RHS vanishes if we assume** $E_1 = E_2$ **or** $\beta_1 = \beta_2$

in all cases $$\boxed{\Delta\phi_{12} = \frac{m_1^2 - m_2^2}{2\mathrm{p}}L = \frac{\Delta m^2}{2E}\mathbf{L}}$$

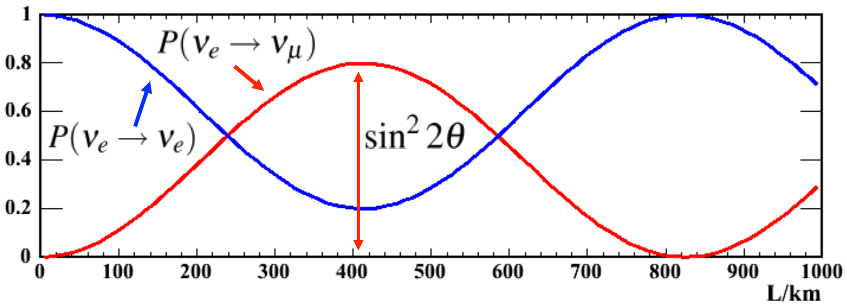
★ **Hence the two-flavour oscillation probability is:**

$$\boxed{P(\nu_e \to \nu_\mu) = \sin^2 2\theta \sin^2\left(\frac{\Delta m_{21}^2 L}{4E}\right)}$$ **with** $$\boxed{\Delta m_{21}^2 = m_2^2 - m_1^2}$$

★ **The corresponding two-flavour survival probability is:**

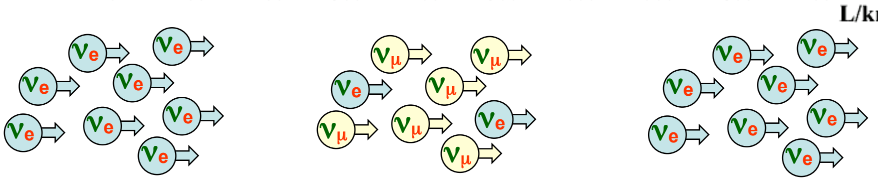
$$\boxed{P(\nu_e \to \nu_e) = 1 - \sin^2 2\theta \sin^2\left(\frac{\Delta m_{21}^2 L}{4E}\right)}$$

•**e.g.** $\Delta m^2 = 0.003\,\mathrm{eV}^2, \quad \sin^2 2\theta = 0.8, \quad E_\nu = 1\,\mathrm{GeV}$



•**wavelength**

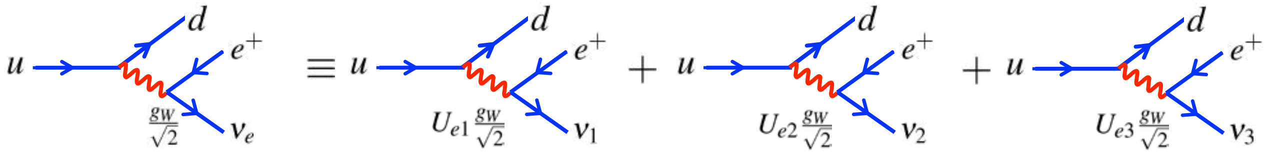
$$\boxed{\lambda_{\mathrm{osc}} = \frac{4\pi E}{\Delta m^2}}$$



Neutrino Oscillations for Three Flavours

★ It is simple to extend this treatment to three generations of neutrinos.
★ In this case we have:

$$\begin{pmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{pmatrix} = \begin{pmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{pmatrix} \begin{pmatrix} \nu_1 \\ \nu_2 \\ \nu_3 \end{pmatrix}$$



★ The 3x3 Unitary matrix U is known as the Pontecorvo-Maki-Nakagawa-Sakata matrix, usually abbreviated PMNS

★ Note : has to be unitary to conserve probability

•Using $\quad U^\dagger U = I \quad \Rightarrow \quad U^{-1} = U^\dagger = (U^*)^T$

gives

$$\begin{pmatrix} \nu_1 \\ \nu_2 \\ \nu_3 \end{pmatrix} = \begin{pmatrix} U^*_{e1} & U^*_{\mu 1} & U^*_{\tau 1} \\ U^*_{e2} & U^*_{\mu 2} & U^*_{\tau 2} \\ U^*_{e3} & U^*_{\mu 3} & U^*_{\tau 3} \end{pmatrix} \begin{pmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{pmatrix}$$

Unitarity Relations

★The Unitarity of the PMNS matrix gives several useful relations: $UU^\dagger = I \quad \Rightarrow$

$$\begin{pmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{pmatrix} \begin{pmatrix} U^*_{e1} & U^*_{\mu 1} & U^*_{\tau 1} \\ U^*_{e2} & U^*_{\mu 2} & U^*_{\tau 2} \\ U^*_{e3} & U^*_{\mu 3} & U^*_{\tau 3} \end{pmatrix} = \begin{pmatrix} 1 & 0 & 0 \\ 0 & 1 & 0 \\ 0 & 0 & 1 \end{pmatrix}$$

gives:

$$U_{e1}U^*_{e1} + U_{e2}U^*_{e2} + U_{e3}U^*_{e3} = 1 \qquad \text{(U1)}$$
$$U_{\mu 1}U^*_{\mu 1} + U_{\mu 2}U^*_{\mu 2} + U_{\mu 3}U^*_{\mu 3} = 1 \qquad \text{(U2)}$$
$$U_{\tau 1}U^*_{\tau 1} + U_{\tau 2}U^*_{\tau 2} + U_{\tau 3}U^*_{\tau 3} = 1 \qquad \text{(U3)}$$
$$U_{e1}U^*_{\mu 1} + U_{e2}U^*_{\mu 2} + U_{e3}U^*_{\mu 3} = 0 \qquad \text{(U4)}$$
$$U_{e1}U^*_{\tau 1} + U_{e2}U^*_{\tau 2} + U_{e3}U^*_{\tau 3} = 0 \qquad \text{(U5)}$$
$$U_{\mu 1}U^*_{\tau 1} + U_{\mu 2}U^*_{\tau 2} + U_{\mu 3}U^*_{\tau 3} = 0 \qquad \text{(U6)}$$

★To calculate the oscillation probability proceed as before…

•Consider a state which is produced at $t = 0$ as a $|\nu_e\rangle$ (i.e. with an electron)

$$|\psi(t=0)\rangle = |\nu_e\rangle = U_{e1}|\nu_1\rangle + U_{e2}|\nu_2\rangle + U_{e3}|\nu_3\rangle$$

•The wave-function evolves as:

$$|\psi(t)\rangle = U_{e1}|\nu_1\rangle e^{-ip_1.x} + U_{e2}|\nu_2\rangle e^{-ip_2.x} + U_{e3}|\nu_3\rangle e^{-ip_3.x}$$

where $p_i.x = E_i t - \vec{p}_i.\vec{x} = E_i t - |\vec{p}|z$

z axis in direction of propagation

•After a travelling a distance L

$$|\psi(L)\rangle = U_{e1}|\nu_1\rangle e^{-i\phi_1} + U_{e2}|\nu_2\rangle e^{-i\phi_2} + U_{e3}|\nu_3\rangle e^{-i\phi_3}$$

where $\phi_i = p_i.x = E_i t - |\vec{p}|L = (E_i - |\vec{p}_i|)L$

•As before we can approximate

$$\phi_i \approx \frac{m_i^2}{2E_i}L$$

•Expressing the mass eigenstates in terms of the weak eigenstates

$$\begin{aligned}|\psi(L)\rangle &= U_{e1}(U_{e1}^*|\nu_e\rangle + U_{\mu 1}^*|\nu_\mu\rangle + U_{\tau 1}^*|\nu_\tau\rangle)e^{-i\phi_1} \\ &+ U_{e2}(U_{e2}^*|\nu_e\rangle + U_{\mu 2}^*|\nu_\mu\rangle + U_{\tau 2}^*|\nu_\tau\rangle)e^{-i\phi_2} \\ &+ U_{e3}(U_{e3}^*|\nu_e\rangle + U_{\mu 3}^*|\nu_\mu\rangle + U_{\tau 3}^*|\nu_\tau\rangle)e^{-i\phi_3}\end{aligned}$$

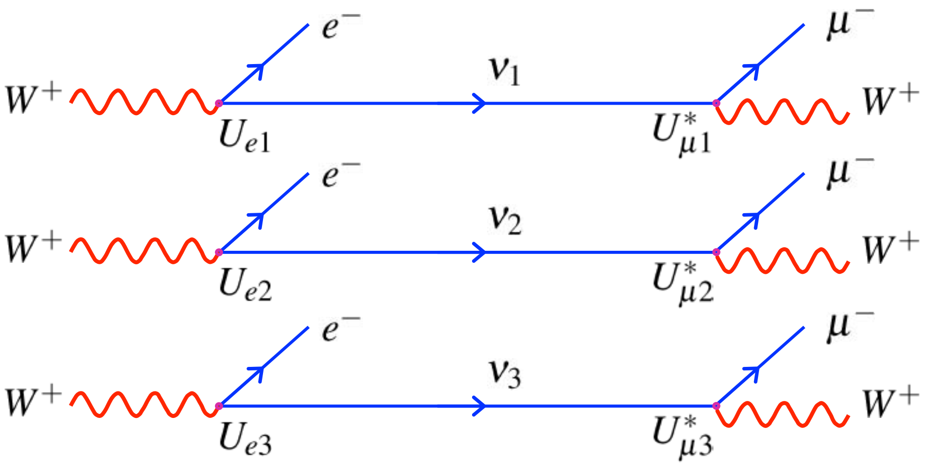
•Which can be rearranged to give

$$\begin{aligned}|\psi(L)\rangle &= (U_{e1}U_{e1}^*e^{-i\phi_1} + U_{e2}U_{e2}^*e^{-i\phi_2} + U_{e3}U_{e3}^*e^{-i\phi_3})|\nu_e\rangle \\ &+ (U_{e1}U_{\mu 1}^*e^{-i\phi_1} + U_{e2}U_{\mu 2}^*e^{-i\phi_2} + U_{e3}U_{\mu 3}^*e^{-i\phi_3})|\nu_\mu\rangle \qquad (3) \\ &+ (U_{e1}U_{\tau 1}^*e^{-i\phi_1} + U_{e2}U_{\tau 2}^*e^{-i\phi_2} + U_{e3}U_{\tau 3}^*e^{-i\phi_3})|\nu_\tau\rangle\end{aligned}$$

•From which

$$\begin{aligned}P(\nu_e \to \nu_\mu) &= |\langle \nu_\mu|\psi(L)\rangle|^2 \\ &= |U_{e1}U_{\mu 1}^*e^{-i\phi_1} + U_{e2}U_{\mu 2}^*e^{-i\phi_2} + U_{e3}U_{\mu 3}^*e^{-i\phi_3}|^2\end{aligned}$$

•The terms in this expression can be represented as:



•Because of the unitarity of the PMNS matrix we have (U4):

$$U_{e1}U_{\mu 1}^* + U_{e2}U_{\mu 2}^* + U_{e3}U_{\mu 3}^* = 0$$

and, consequently, unless the phases of the different components are different, the sum of these three diagrams is zero, i.e., require different neutrino masses for osc.

• **Evaluate**

$$P(\nu_e \to \nu_\mu) = |U_{e1}U^*_{\mu 1}e^{-i\phi_1} + U_{e2}U^*_{\mu 2}e^{-i\phi_2} + U_{e3}U^*_{\mu 3}e^{-i\phi_3}|^2$$

using $|z_1+z_2+z_3|^2 \equiv |z_1|^2+|z_2|^2+|z_3|^2+2\Re(z_1z_2^*+z_1z_3^*+z_2z_3^*)$ **(4)**

which gives:

$$\begin{aligned} P(\nu_e \to \nu_\mu) &= |U_{e1}U^*_{\mu 1}|^2 + |U_{e2}U^*_{\mu 2}|^2 + |U_{e3}U^*_{\mu 3}|^2 + \\ &2\Re(U_{e1}U^*_{\mu 1}U^*_{e2}U_{\mu 2}e^{-i(\phi_1-\phi_2)} + U_{e1}U^*_{\mu 1}U^*_{e3}U_{\mu 3}e^{-i(\phi_1-\phi_3)} + U_{e2}U^*_{\mu 2}U^*_{e3}U_{\mu 3}e^{-i(\phi_2-\phi_3)}) \end{aligned} \quad \textbf{(5)}$$

• **This can be simplified by applying identity (4) to |(U4)|²**

$$|U_{e1}U^*_{\mu 1} + U_{e2}U^*_{\mu 2} + U_{e3}U^*_{\mu 3}|^2 = 0$$

$$\Rightarrow \quad |U_{e1}U^*_{\mu 1}|^2 + |U_{e2}U^*_{\mu 2}|^2 + |U_{e3}U^*_{\mu 3}|^2 = -2\Re(U_{e1}U^*_{\mu 1}U^*_{e2}U_{\mu 2} + U_{e1}U^*_{\mu 1}U^*_{e3}U_{\mu 3} + U_{e2}U^*_{\mu 2}U^*_{e3}U_{\mu 3})$$

• **Substituting into equation (5) gives**

$$\boxed{\begin{aligned} P(\nu_e \to \nu_\mu) &= 2\Re\{U_{e1}U^*_{\mu 1}U^*_{e2}U_{\mu 2}[e^{-i(\phi_1-\phi_2)}-1]\} \\ &+ 2\Re\{U_{e1}U^*_{\mu 1}U^*_{e3}U_{\mu 3}[e^{-i(\phi_1-\phi_3)}-1]\} \\ &+ 2\Re\{U_{e2}U^*_{\mu 2}U^*_{e3}U_{\mu 3}[e^{-i(\phi_2-\phi_3)}-1]\} \end{aligned}} \quad \textbf{(6)}$$

★ **This expression for the electron survival probability is obtained from the coefficient for $|\nu_e\rangle$ in eqn. (3):**

$$\begin{aligned} P(\nu_e \to \nu_e) &= |\langle \nu_e|\psi(L)\rangle|^2 \\ &= |U_{e1}U^*_{e1}e^{-i\phi_1} + U_{e2}U^*_{e2}e^{-i\phi_2} + U_{e3}U^*_{e3}e^{-i\phi_3}|^2 \end{aligned}$$

which using the unitarity relation (U1)

$$|U_{e1}U^*_{e1} + U_{e2}U^*_{e2} + U_{e3}U^*_{e3}|^2 = 1$$

can be written

$$\boxed{\begin{aligned} P(\nu_e \to \nu_e) = 1 &+ 2|U_{e1}|^2|U_{e2}|^2\Re\{[e^{-i(\phi_1-\phi_2)}-1]\} \\ &+ 2|U_{e1}|^2|U_{e3}|^2\Re\{[e^{-i(\phi_1-\phi_3)}-1]\} \\ &+ 2|U_{e2}|^2|U_{e3}|^2\Re\{[e^{-i(\phi_2-\phi_3)}-1]\} \end{aligned}} \quad \textbf{(7)}$$

★ **This expression can simplified using**

$$\begin{aligned} \Re\{e^{-i(\phi_1-\phi_2)}-1\} &= \cos(\phi_2-\phi_1)-1 \\ &= -2\sin^2\left(\frac{\phi_2-\phi_1}{2}\right) \\ &= -2\sin^2\left(\frac{(m_2^2-m_1^2)L}{4E}\right) \end{aligned}$$

with $\phi_i \approx \frac{m_i^2}{2E}L$

Phase of mass eigenstate i at $z = L$

•Define: $\boxed{\Delta_{21} = \frac{(m_2^2 - m_1^2)L}{4E} = \frac{\Delta m_{21}^2 L}{4E}}$ with $\boxed{\Delta m_{21}^2 = m_2^2 - m_1^2}$

NOTE: $\Delta_{21} = (\phi_2 - \phi_1)/2$ **is a phase difference (i.e. dimensionless)**

•Which gives the electron neutrino survival probability

$$\boxed{P(\nu_e \to \nu_e) = 1 - 4|U_{e1}|^2|U_{e2}|^2 \sin^2\Delta_{21} - 4|U_{e1}|^2|U_{e3}|^2 \sin^2\Delta_{31} - 4|U_{e2}|^2|U_{e3}|^2 \sin^2\Delta_{32}}$$

•Similar expressions can be obtained for the muon and tau neutrino survival probabilities for muon and tau neutrinos.

★ Note that since we only have three neutrino generations there are only two independent mass-squared differences, i.e.

$$m_3^2 - m_1^2 = (m_3^2 - m_2^2) + (m_2^2 - m_1^2)$$

and in the above equation only two of the Δ_{ij} are independent

★All expressions are in Natural Units, conversion to more useful units here gives:

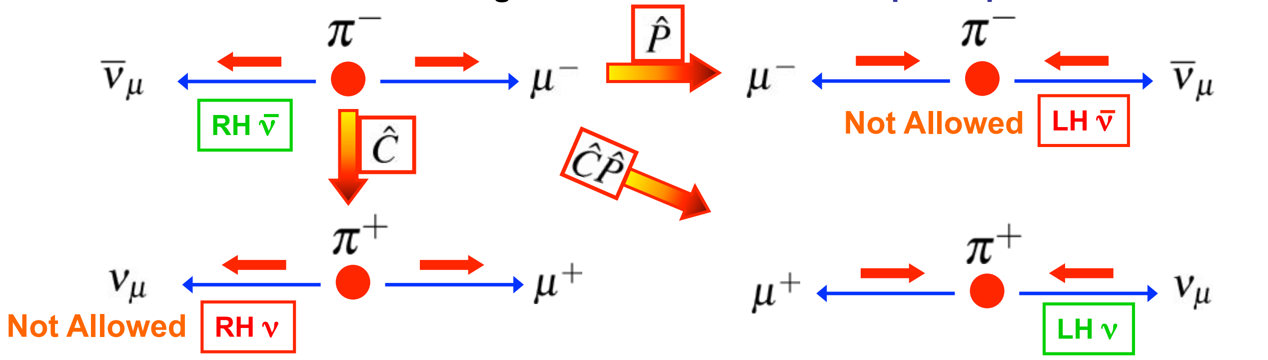
$$\boxed{\Delta_{21} = 1.27\frac{\Delta m_{21}^2(\mathrm{eV}^2)L(\mathrm{km})}{E(\mathrm{GeV})}} \quad \text{and} \quad \boxed{\lambda_{\mathrm{osc}}(\mathrm{km}) = 2.47\frac{E(\mathrm{GeV})}{\Delta m^2(\mathrm{eV}^2)}}$$

CP and CPT in the Weak Interaction

★ In addition to parity there are two other important discrete symmetries:

Parity	$\hat{P}:$	$\vec{r} \to -\vec{r}$
Time Reversal	$\hat{T}:$	$t \to -t$
Charge Conjugation	$\hat{C}:$	Particle ⟷ Anti-particle

★ The weak interaction violates parity conservation, but what about **C** ? Consider pion decay remembering that the neutrino is ultra-relativistic and only **left-handed neutrinos** and **right-handed anti-neutrinos** participate in WI



★ Hence weak interaction also **violates charge conjugation** symmetry but appears to be invariant under combined effect of **C** and **P**

CP transforms:

RH Particles ⟺ LH Anti-particles

LH Particles ⟺ RH Anti-particles

★ If the weak interaction were invariant under CP expect

$$\boxed{\Gamma(\pi^+ \to \mu^+ \nu_\mu) = \Gamma(\pi^- \to \mu^- \overline{\nu}_\mu)}$$

★ All Lorentz invariant Quantum Field Theories can be shown to be invariant under CPT (charge conjugation + parity + time reversal)

➡ Particles/anti-particles have identical mass, lifetime, magnetic moments,…

Best current experimental test: $m_{K^0} - m_{\overline{K}^0} < 6 \times 10^{-19} m_{K^0}$

★ Believe CPT has to hold:

if CP invariance holds ➡ time reversal symmetry

if CP is violated ➡ time reversal symmetry violated

★ To account for the small excess of matter over anti-matter that must have existed early in the universe require CP violation in particle physics !

★ CP violation can arise in the weak interaction (see also handout 12).

CP and T Violation in Neutrino Oscillations

• Previously derived the oscillation probability for $\nu_e \to \nu_\mu$

$$\begin{aligned} P(\nu_e \to \nu_\mu) &= 2\Re\{U_{e1}U_{\mu1}^*U_{e2}^*U_{\mu2}[e^{-i(\phi_1-\phi_2)}-1]\} \\ &+ 2\Re\{U_{e1}U_{\mu1}^*U_{e3}^*U_{\mu3}[e^{-i(\phi_1-\phi_3)}-1]\} \\ &+ 2\Re\{U_{e2}U_{\mu2}^*U_{e3}^*U_{\mu3}[e^{-i(\phi_2-\phi_3)}-1]\} \end{aligned}$$

• The oscillation probability for $\nu_\mu \to \nu_e$ can be obtained in the same manner or by simply exchanging the labels $(e) \leftrightarrow (\mu)$

$$\begin{aligned} P(\nu_\mu \to \nu_e) &= 2\Re\{U_{\mu1}U_{e1}^*U_{\mu2}^*U_{e2}[e^{-i(\phi_1-\phi_2)}-1]\} \\ &+ 2\Re\{U_{\mu1}U_{e1}^*U_{\mu3}^*U_{e3}[e^{-i(\phi_1-\phi_3)}-1]\} \\ &+ 2\Re\{U_{\mu2}U_{e2}^*U_{\mu3}^*U_{e3}[e^{-i(\phi_2-\phi_3)}-1]\} \end{aligned} \quad (8)$$

★ Unless the elements of the PMNS matrix are real (see note below)

$$P(\nu_e \to \nu_\mu) \neq P(\nu_\mu \to \nu_e) \quad (9)$$

• If any of the elements of the PMNS matrix are complex, neutrino oscillations are not invariant under time reversal $t \to -t$

NOTE: can multiply entire PMNS matrix by a complex phase without changing the oscillation prob. T is violated if one of the elements has a different complex phase than the others

•Consider the effects of T, CP and CPT on neutrino oscillations

$\boxed{\text{T}}\quad \nu_e \to \nu_\mu \xrightarrow{\hat{T}} \nu_\mu \to \nu_e$

$\boxed{\text{CP}}\quad \nu_e \to \nu_\mu \xrightarrow{\hat{C}\hat{P}} \overline{\nu}_e \to \overline{\nu}_\mu$

$\boxed{\text{CPT}}\quad \nu_e \to \nu_\mu \xrightarrow{\hat{C}\hat{P}\hat{T}} \overline{\nu}_\mu \to \overline{\nu}_e$

Note C alone is not sufficient as it transforms LH neutrinos into LH anti-neutrinos (not involved in Weak Interaction)

•If the weak interactions is invariant under CPT

$$P(\nu_e \to \nu_\mu) = P(\overline{\nu}_\mu \to \overline{\nu}_e)$$

and similarly
$$P(\nu_\mu \to \nu_e) = P(\overline{\nu}_e \to \overline{\nu}_\mu) \qquad (10)$$

•If the PMNS matrix is not purely real, then (9)

$$P(\nu_e \to \nu_\mu) \neq P(\nu_\mu \to \nu_e)$$

and from (10)
$$\boxed{P(\nu_e \to \nu_\mu) \neq P(\overline{\nu}_e \to \overline{\nu}_\mu)}$$

★Hence unless the PMNS matrix is real, **CP is violated** in neutrino oscillations!

Future experiments, e.g. "a neutrino factory", are being considered as a way to investigate CP violation in neutrino oscillations. However, CP violating effects are well below the current experimental sensitivity. In the following discussion we will take the PMNS matrix to be real. (question 22)

Neutrino Mass Hierarchy

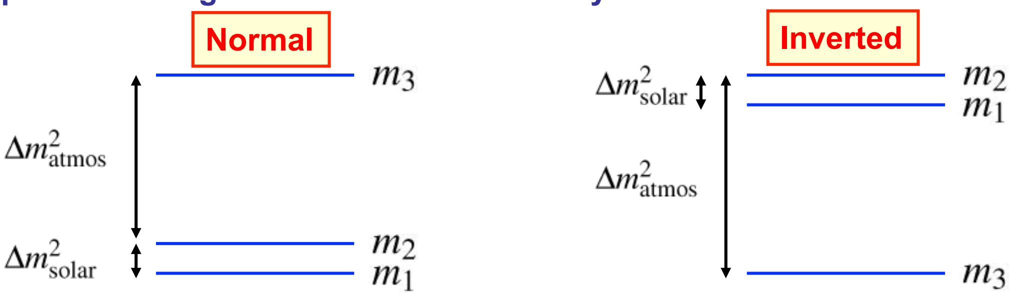
★ To date, results on neutrino oscillations only determine

$$|\Delta m^2_{ji}| = |m^2_j - m^2_i|$$

★ Two distinct and very different mass scales:

- Atmospheric neutrino oscillations : $|\Delta m^2|_{\text{atmos}} \sim 2.5 \times 10^{-3}\,\text{eV}^2$
- Solar neutrino oscillations: $|\Delta m^2|_{\text{solar}} \sim 8 \times 10^{-5}\,\text{eV}^2$

•Two possible assignments of mass hierarchy:



•In both cases: $\Delta m^2_{21} \sim 8 \times 10^{-5}\,\text{eV}^2$ (solar)

$$|\Delta m^2_{31}| \approx |\Delta m^2_{32}| \sim 2.5 \times 10^{-3}\,\text{eV}^2 \quad \text{(atmospheric)}$$

•Hence we can approximate $\Delta m^2_{31} \approx \Delta m^2_{32}$

Three Flavour Oscillations Neglecting CP Violation

- Neglecting CP violation considerably simplifies the algebra of three flavour neutrino oscillations. Taking the PMNS matrix to be real, equation (6) becomes:

$$P(\nu_e \to \nu_\mu) = -4U_{e1}U_{\mu 1}U_{e2}U_{\mu 2}\sin^2\Delta_{21} - 4U_{e1}U_{\mu 1}U_{e3}U_{\mu 3}\sin^2\Delta_{31} - 4U_{e2}U_{\mu 2}U_{e3}U_{\mu 3}\sin^2\Delta_{32}$$

with $\Delta_{ji} = \dfrac{(m_j^2 - m_i^2)L}{4E} = \dfrac{\Delta m_{ji}^2 L}{4E}$

- Using: $\Delta_{31} \approx \Delta_{32}$ (see p. 365)

$$P(\nu_e \to \nu_\mu) \approx -4U_{e1}U_{\mu 1}U_{e2}U_{\mu 2}\sin^2\Delta_{21} - 4(U_{e1}U_{\mu 1} + U_{e2}U_{\mu 2})U_{e3}U_{\mu 3}\sin^2\Delta_{32}$$

- Which can be simplified using (U4) $U_{e1}U_{\mu 1}^* + U_{e2}U_{\mu 2}^* + U_{e3}U_{\mu 3}^* = 0$

$$\Rightarrow \boxed{P(\nu_e \to \nu_\mu) \approx -4U_{e1}U_{\mu 1}U_{e2}U_{\mu 2}\sin^2\Delta_{21} + 4U_{e3}^2U_{\mu 3}^2\sin^2\Delta_{32}}$$

- Can apply $\Delta_{31} \approx \Delta_{32}$ to the expression for electron neutrino survival probability

$$P(\nu_e \to \nu_e) = 1 - 4U_{e1}^2U_{e2}^2\sin^2\Delta_{21} - 4U_{e1}^2U_{e3}^2\sin^2\Delta_{31} - 4U_{e2}^2U_{e3}^2\sin^2\Delta_{32}$$
$$\approx 1 - 4U_{e1}^2U_{e2}^2\sin^2\Delta_{21} - 4(U_{e1}^2 + U_{e2}^2)U_{e3}^2\sin^2\Delta_{32}$$

- Which can be simplified using (U1) $U_{e1}^2 + U_{e2}^2 + U_{e3}^2 = 1$

$$\Rightarrow \boxed{P(\nu_e \to \nu_e) \approx 1 - 4U_{e1}^2U_{e2}^2\sin^2\Delta_{21} - 4(1 - U_{e3}^2)U_{e3}^2\sin^2\Delta_{32}}$$

★ Neglecting CP violation (i.e. taking the PMNS matrix to be real) and making the approximation that $|\Delta m_{31}^2| \approx |\Delta m_{32}^2|$ obtain the following expressions for neutrino oscillation probabilities:

$$P(\nu_e \to \nu_e) \approx 1 - 4U_{e1}^2U_{e2}^2\sin^2\Delta_{21} - 4(1 - U_{e3}^2)U_{e3}^2\sin^2\Delta_{32} \quad (11)$$

$$P(\nu_\mu \to \nu_\mu) \approx 1 - 4U_{\mu 1}^2U_{\mu 2}^2\sin^2\Delta_{21} - 4(1 - U_{\mu 3}^2)U_{\mu 3}^2\sin^2\Delta_{32} \quad (12)$$

$$P(\nu_\tau \to \nu_\tau) \approx 1 - 4U_{\tau 1}^2U_{\tau 2}^2\sin^2\Delta_{21} - 4(1 - U_{\tau 3}^2)U_{\tau 3}^2\sin^2\Delta_{32} \quad (13)$$

$$P(\nu_e \to \nu_\mu) = P(\nu_\mu \to \nu_e) \approx -4U_{e1}U_{\mu 1}U_{e2}U_{\mu 2}\sin^2\Delta_{21} + 4U_{e3}^2U_{\mu 3}^2\sin^2\Delta_{32} \quad (14)$$

$$P(\nu_e \to \nu_\tau) = P(\nu_\tau \to \nu_e) \approx -4U_{e1}U_{\tau 1}U_{e2}U_{\tau 2}\sin^2\Delta_{21} + 4U_{e3}^2U_{\tau 3}^2\sin^2\Delta_{32} \quad (15)$$

$$P(\nu_\mu \to \nu_\tau) = P(\nu_\tau \to \nu_\mu) \approx -4U_{\mu 1}U_{\tau 1}U_{\mu 2}U_{\tau 2}\sin^2\Delta_{21} + 4U_{\mu 3}^2U_{\tau 3}^2\sin^2\Delta_{32} \quad (16)$$

★ The wavelengths associated with $\sin^2\Delta_{21}$ and $\sin^2\Delta_{32}$ are:

"SOLAR" $\lambda_{21} = \dfrac{4\pi E}{\Delta m_{21}^2}$ "Long"-Wavelength

and

$\lambda_{32} = \dfrac{4\pi E}{\Delta m_{32}^2}$ "ATMOSPHERIC" "Short"-Wavelength

PMNS Matrix

★ The PMNS matrix is usually expressed in terms of 3 rotation angles $\theta_{12}, \theta_{23}, \theta_{13}$ and a complex phase δ, using the notation $s_{ij} = \sin\theta_{ij}, \quad c_{ij} = \cos\theta_{ij}$

$$\begin{pmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{pmatrix} = \underbrace{\begin{pmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{pmatrix}}_{\text{"Atmospheric"}} \times \begin{pmatrix} c_{13} & 0 & s_{13}e^{-i\delta} \\ 0 & 1 & 0 \\ -s_{13}e^{-i\delta} & 0 & c_{13} \end{pmatrix} \times \underbrace{\begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix}}_{\text{"Solar"}}$$

Dominates: "Atmospheric" "Solar"

• Writing this out in full:

$$U = \begin{pmatrix} c_{12}c_{13} & s_{12}c_{13} & s_{13}e^{-i\delta} \\ -s_{12}c_{23} - c_{12}s_{23}s_{13}e^{i\delta} & c_{12}c_{23} - s_{12}s_{23}s_{13}e^{i\delta} & s_{23}c_{13} \\ s_{12}s_{23} - c_{12}c_{23}s_{13}e^{i\delta} & -c_{12}s_{23} - s_{12}c_{23}s_{13}e^{i\delta} & c_{23}c_{13} \end{pmatrix}$$

★There are six <u>SM parameters</u> that can be measured in ν oscillation experiments

$\lvert\Delta m_{21}\rvert^2 = \lvert m_2^2 - m_1^2\rvert$	θ_{12}	Solar and reactor neutrino experiments
$\lvert\Delta m_{32}\rvert^2 = \lvert m_3^2 - m_2^2\rvert$	θ_{23}	Atmospheric and beam neutrino experiments
	θ_{13}	Reactor neutrino experiments + future beam
	δ	Future beam experiments

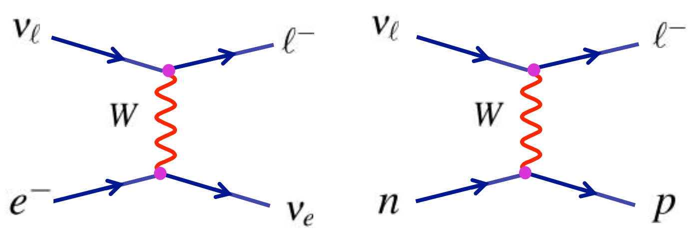
Neutrino Experiments

•Before discussing current experimental data, need to consider how neutrinos interact in matter (i.e. our detectors)

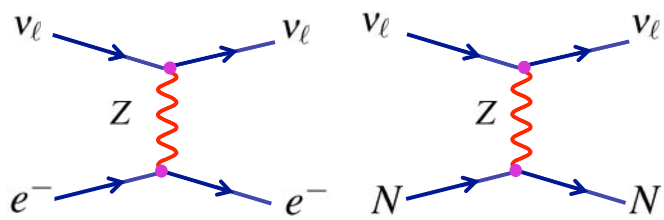
<u>Two processes:</u>
- Charged current (CC) interactions (via a W-boson) ⇨ charged lepton
- Neutral current (NC) interactions (via a Z-boson)

<u>Two possible "targets":</u> can have neutrino interactions with
- atomic electrons
- nucleons within the nucleus

CHARGED CURRENT



NEUTRAL CURRENT



Neutrino Interaction Thresholds

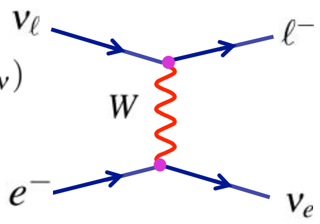
★ Neutrino detection method depends on the neutrino energy and (weak) flavour

- Neutrinos from the sun and nuclear reactions have $E_\nu \sim 1\,\mathrm{MeV}$
- Atmospheric neutrinos have $E_\nu \sim 1\,\mathrm{GeV}$

★ These energies are relatively low and not all interactions are kinematically allowed, i.e. there is a threshold energy before an interaction can occur. Require sufficient energy in the centre-of-mass frame to produce the final state particles

❶ Charged current interactions on atomic electrons (in laboratory frame)

$$p_\nu = (E_\nu, 0, 0, E_\nu)$$
$$p_e = (m_e, 0, 0, 0)$$



$$s = (p_\nu + p_e)^2 = (E_\nu + m_e)^2 - E_\nu^2$$

Require: $s > m_\ell^2$

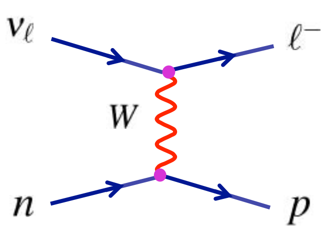
$$\Rightarrow \boxed{E_\nu > \left[\left(\frac{m_\ell}{m_e}\right)^2 - 1\right]\frac{m_e}{2}}$$

- Putting in the numbers, for CC interactions with atomic electrons require

$$E_{\nu_e} > 0 \qquad \underbrace{E_{\nu_\mu} > 11\,\mathrm{GeV} \qquad E_{\nu_\tau} > 3090\,\mathrm{GeV}}$$

High energy thresholds compared to typical energies considered here

❷ charged current interactions on nucleons (in lab. frame)



$$s = (p_\nu + p_n)^2 = (E_\nu + m_n)^2 - E_\nu^2$$

Require: $s > (m_\ell + m_p)^2$

$$\Rightarrow \boxed{E_\nu > \frac{(m_p^2 - m_n^2) + m_\ell^2 + 2m_p m_\ell}{2m_n}}$$

- For CC interactions from neutrons require

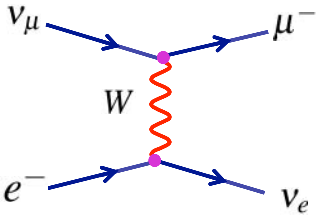
$$E_{\nu_e} > 0 \qquad E_{\nu_\mu} > 110\,\mathrm{MeV} \qquad E_{\nu_\tau} > 3.5\,\mathrm{GeV}$$

★ Electron neutrinos from the sun and nuclear reactors $E_\nu \sim 1\,\mathrm{MeV}$ which oscillate into muon or tau neutrinos cannot interact via charged current interactions – "they effectively disappear"

★ Atmospheric muon neutrinos $E_\nu \sim 1\,\mathrm{GeV}$ which oscillate into tau neutrinos cannot interact via charged current interactions – "disappear"

- To date, most experimental signatures for neutrino oscillation are a deficit of neutrino interactions (with the exception of SNO) because below threshold for produce lepton of different flavour from original neutrino

- In Handout 10 derived expressions for CC neutrino-quark cross sections in ultra-relativistic limit (neglecting masses of neutrinos/quarks)
- For **high energy muon** neutrinos can directly use the results from page 316



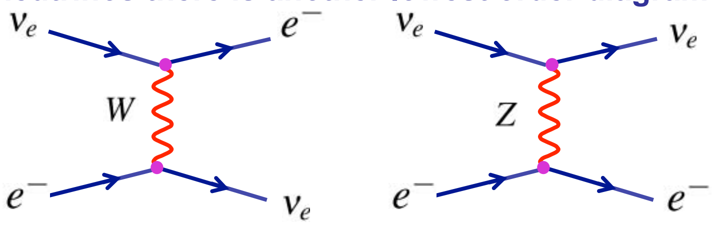
$$\sigma_{\nu_\mu e^-} = \frac{G_F^2 s}{\pi}$$

with $\quad s = (E_\nu + m_e)^2 - E_\nu^2 \approx 2m_e E_\nu$

$$\boxed{\sigma_{\nu_\mu e^-} = \frac{2m_e G_F^2 E_\nu}{\pi}}$$

Cross section increases linearly with lab. frame neutrino energy

- For **electron** neutrinos there is another lowest order diagram with the same final state



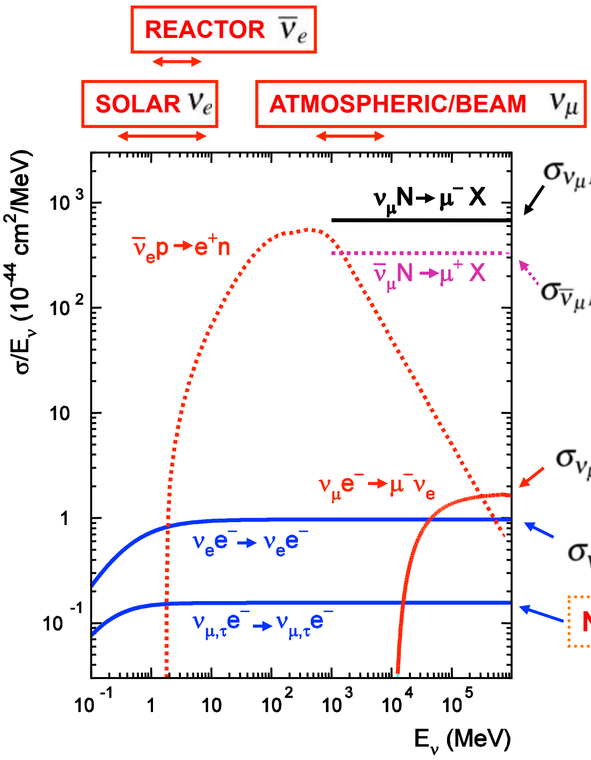
It turns out that the cross section is lower than the pure CC cross section due to negative interference when summing matrix elements $|M_{CC} + M_{NC}|^2 < |M_{CC}|^2$

$$\boxed{\sigma_{\nu_e e} \approx 0.6\sigma_{\nu_e e}^{CC}}$$

- In the high energy limit the CC neutrino-nucleon cross sections are larger due to the higher centre-of-mass energy: $s = (E_\nu + m_n)^2 - E_\nu^2 \approx 2m_n E_\nu$

Neutrino Detection

★ The detector technology/interaction process depends on type of neutrino and energy



$$\sigma_{\nu_\mu N} = 0.22\frac{2m_p G_F^2 E_\nu}{\pi}$$

$$\sigma_{\overline{\nu}_\mu N} = 0.33\frac{2m_p G_F^2 E_\nu}{3\pi}$$

Deep Inelastic Scattering (p. 332)

$$\sigma_{\nu_\mu e^-} = \frac{2m_e G_F^2 E_\nu}{\pi}$$

CC only
Threshold 11 GeV

$$\sigma_{\nu_e e^-} \approx 0.6\frac{2m_e G_F^2 E_\nu}{\pi}$$

NC + CC

NC only (see handout 13)

Atmospheric/Beam Neutrinos $\nu_e, \nu_\mu, \overline{\nu}_e, \overline{\nu}_\mu : E_\nu > 1\,\mathrm{GeV}$

❶ <u>Water Čerenkov:</u> e.g. Super Kamiokande

❷ <u>Iron Calorimeters:</u> e.g. MINOS, CDHS (see handout 10)

- Produce high energy charged lepton – relatively easy to detect

Solar Neutrinos $\nu_e \;:\; E_\nu < 20\,\mathrm{MeV}$

❶ <u>Water Čerenkov:</u> e.g. Super Kamiokande

- Detect Čerenkov light from electron produced in $\nu_e + e^- \rightarrow \nu_e + e^-$
- Because of background from natural radioactivity limited to $E_\nu > 5\,\mathrm{MeV}$
- Because Oxygen is a doubly magic nucleus don't get $\nu_e + n \rightarrow e^- + p$

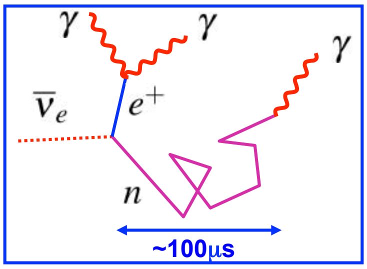
❷ <u>Radio-Chemical:</u> e.g. Homestake, SAGE, GALLEX

- Use inverse beta decay process, e.g. $\nu_e + {}^{71}\mathrm{Ga} \rightarrow e^- + {}^{71}\mathrm{Ge}$
- Chemically extract produced isotope and count decays (only gives a rate)

Reactor Neutrinos $\overline{\nu}_e \;:\; E_{\overline{\nu}} < 5\,\mathrm{MeV}$

❶ <u>Liquid Scintillator:</u> e.g. KamLAND

- Low energies ➡ large radioactive background
- Dominant interaction: $\overline{\nu}_e + p \rightarrow e^+ + n$
- Prompt positron annihilation signal + delayed signal from n (space/time correlation reduces background)
- electrons produced by photons excite scintillator which produces light



1) Long Baseline Neutrino Experiments

- Initial studies of neutrino oscillations from atmospheric and solar neutrinos
  - atmospheric neutrinos discussed in **examinable** appendix
- Emphasis of neutrino research now on neutrino beam experiments
- Allows the physicist to take control – design experiment with specific goals
- In the last few years, long baseline neutrino oscillation experiments have started taking data: K2K, MINOS, CNGS, T2K

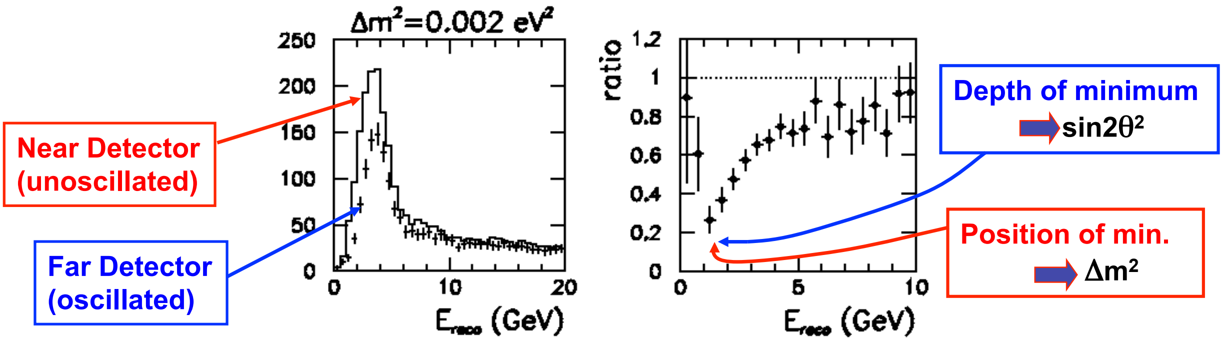
<u>**Basic Idea:**</u>

★ Intense neutrino beam

★ Two detectors: one close to beam the other hundreds of km away

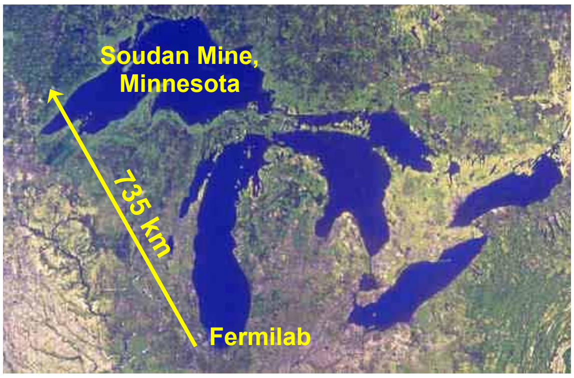
★ Measure ratio of the neutrino energy spectrum in far detector (oscillated) to that in the near detector (unoscillated)

★ Partial cancellation of systematic biases

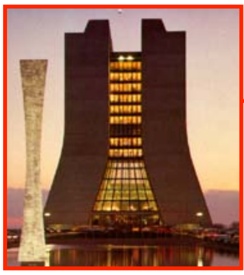


MINOS

- 120 GeV protons extracted from the MAIN INJECTOR at Fermilab (see p. 271)
- 2.5×10^{13} protons per pulse hit target ➡ very intense beam - 0.3 MW on target



Two detectors:

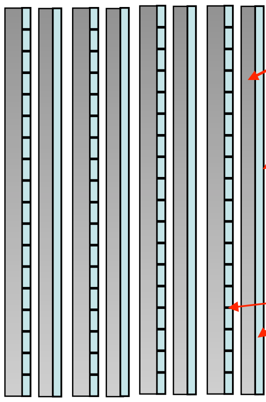


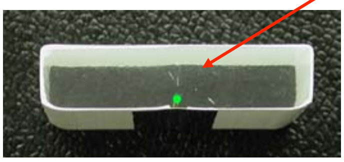
★ 1000 ton, NEAR Detector at Fermilab : 1 km from beam

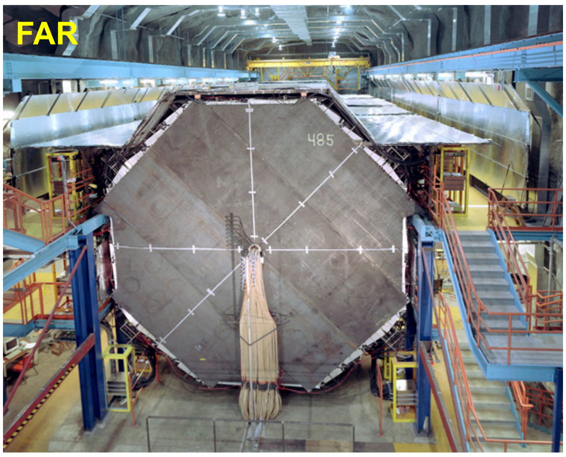
★ 5400 ton FAR Detector, 720m underground in Soudan mine, N. Minnesota: 735 km from beam

The MINOS Detectors:

- Dealing with high energy neutrinos $E_\nu > 1\,\mathrm{GeV}$
- The muons produced by ν_μ interactions travel several metres
- Steel-Scintillator sampling calorimeter
- Each plane: 2.54 cm steel +1 cm scintillator
- A charged particle crossing the scintillator produces light – detect with PMTs

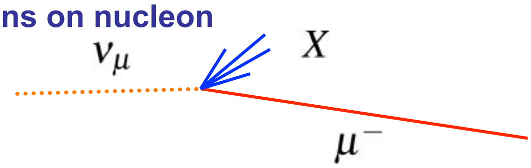






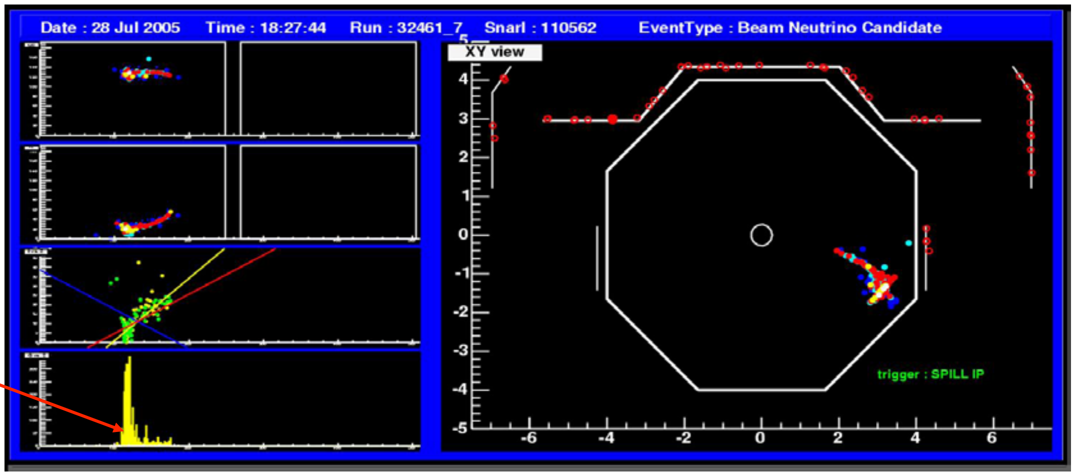
•Neutrino detection via CC interactions on nucleon

$$\nu_\mu + N \rightarrow \mu^- + X$$



Example event:

Signal from hadronic shower



•The main feature of the MINOS detector is the very good neutrino energy resolution

$$\boxed{E_\nu = E_\mu + E_X}$$

•Muon energy from range/curvature in B-field
•Hadronic energy from amount of light observed

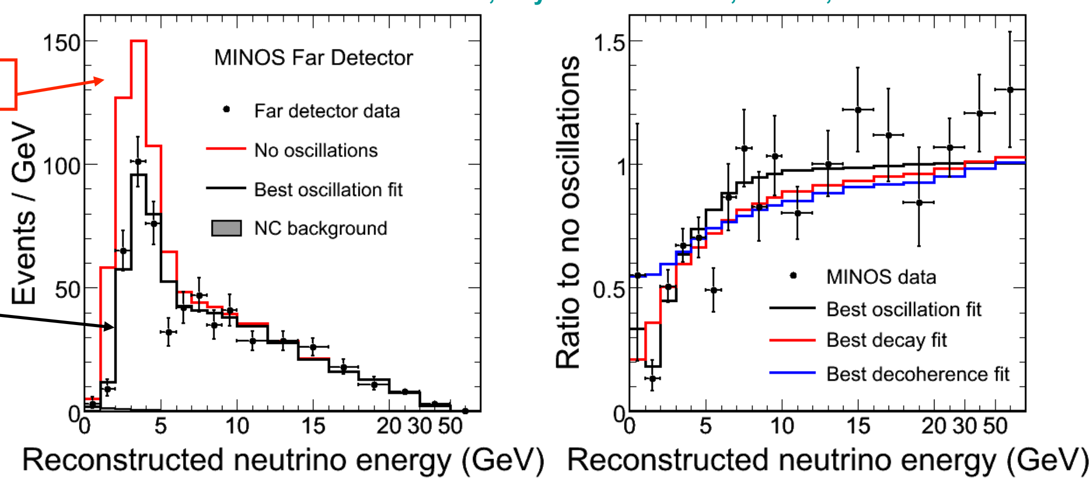
MINOS Results

- For the MINOS experiment L is fixed and observe oscillations as function of E_ν
- For $|\Delta m^2_{32}| \sim 2.5 \times 10^{-3}\,\mathrm{eV}^2$ first oscillation minimum at $E_\nu = 1.5\,\mathrm{GeV}$
- To a very good approximation can use two flavour formula as oscillations corresponding to $|\Delta m^2_{21}| \sim 8 \times 10^{-5}\,\mathrm{eV}^2$ occur at $E_\nu = 50\,\mathrm{MeV}$, beam contains very few neutrinos at this energy + well below detection threshold

MINOS Collaboration, Phys. Rev. Lett. 101, 131802, 2008

no oscillations

best fit osc

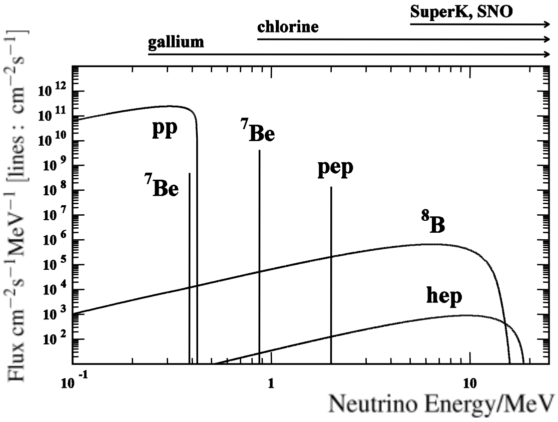


$$\boxed{|\Delta m^2_{32}| = (2.43 \pm 0.12) \times 10^{-3}\,\mathrm{eV}^2}$$

2) Solar Neutrinos

•The Sun is powered by the weak interaction – producing a very large flux of electron neutrinos

$$\boxed{2 \times 10^{38} \nu_e \, s^{-1}}$$



•Several different nuclear reactions in the sun ⇨complex neutrino energy spectrum

$$\boxed{p+p \rightarrow d+e^+ + \nu_e} \quad E_\nu < 0.5\,\mathrm{MeV}$$

$$\boxed{^8B \rightarrow {}^8Be^* + e^+ + \nu_e} \quad E_\nu \sim 5\,\mathrm{MeV}$$

$$p+e^- + p \rightarrow d + \nu_e$$

$$^7Be + e^- \rightarrow {}^7Li + \nu_e$$

$$^3He + p \rightarrow {}^4He + e^+ + \nu_e$$

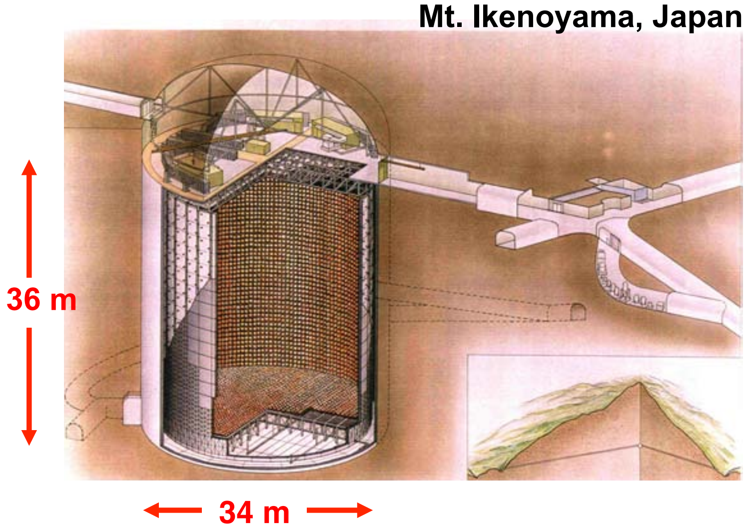
•All experiments saw a deficit of electron neutrinos compared to experimental prediction – the SOLAR NEUTRINO PROBLEM

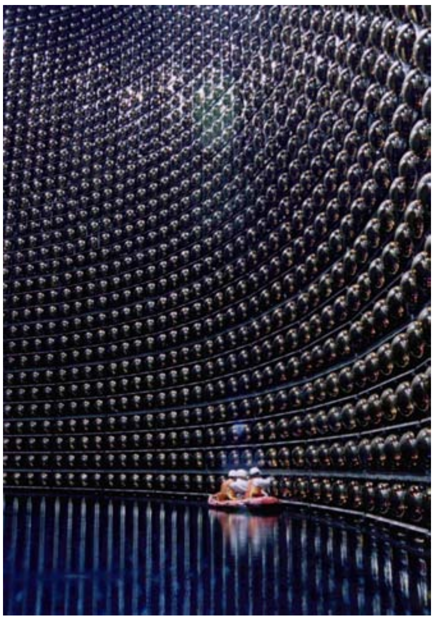
• e.g. Super Kamiokande

Solar Neutrinos I: Super Kamiokande

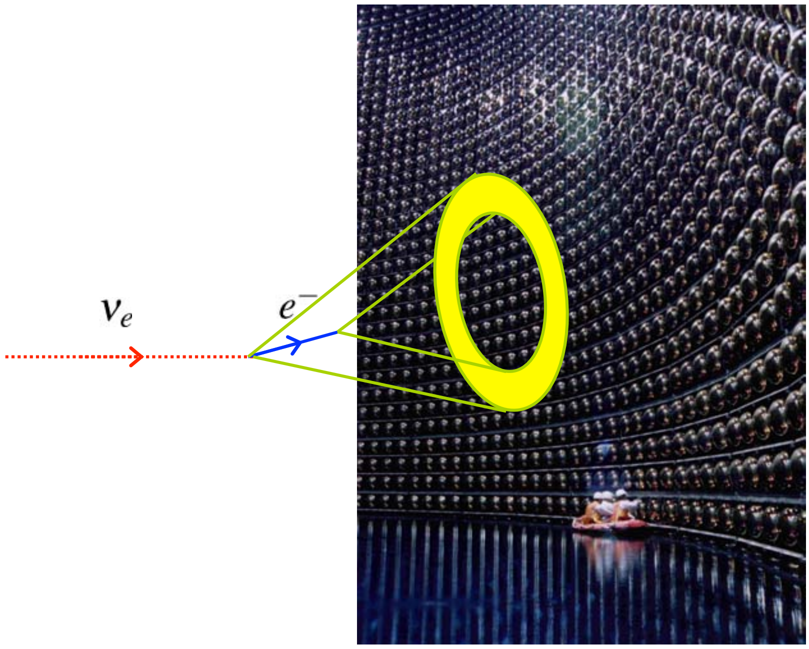
- 50000 ton water Čerenkov detector
- Water viewed by 11146 Photo-multiplier tubes
- Deep underground to filter out cosmic rays otherwise difficult to detect rare neutrino interactions

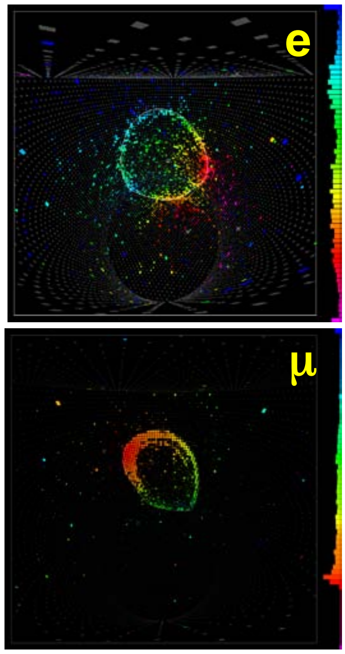
Mt. Ikenoyama, Japan





•Detect neutrinos by observing Čerenkov radiation from charged particles which travel faster than speed of light in water *c/n*





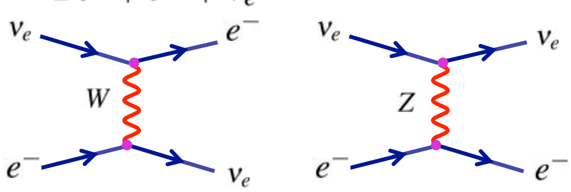
•Can distinguish electrons from muons from pattern of light – muons produce clean rings whereas electrons produce more diffuse "fuzzy" rings

- Sensitive to solar neutrinos with $E_\nu > 5\,\mathrm{MeV}$
- For lower energies too much background from natural radioactivity (β-decays)
- Hence detect mostly neutrinos from $^8B \rightarrow ^8Be^* + e^+ + \nu_e$

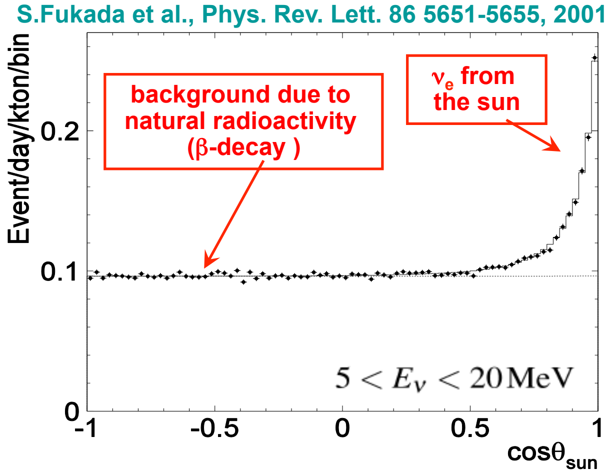
•Detect electron Čerenkov rings from

$$\nu_e + e^- \rightarrow \nu_e + e^-$$

•In LAB frame the electron is produced preferentially along the ν_e direction



S.Fukada et al., Phys. Rev. Lett. 86 5651-5655, 2001



Results:

- Clear signal of neutrinos from the sun
- However too few neutrinos

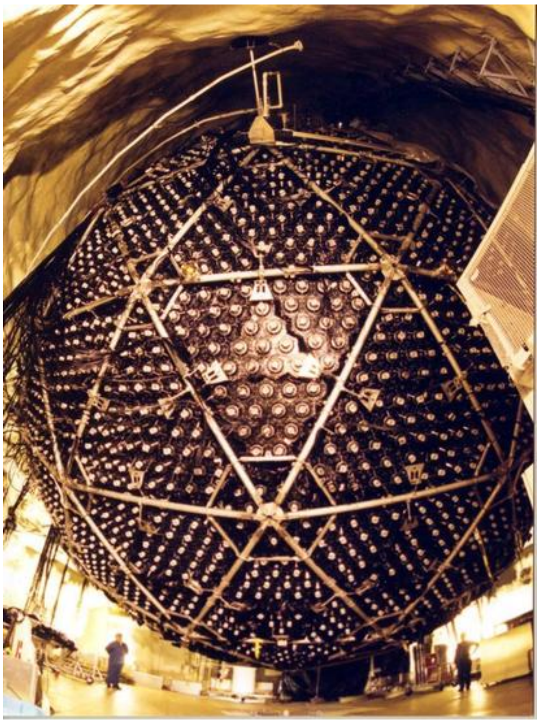
$\boxed{\text{DATA/SSM} = 0.45 \pm 0.02}$

SSM = "Standard Solar Model" Prediction

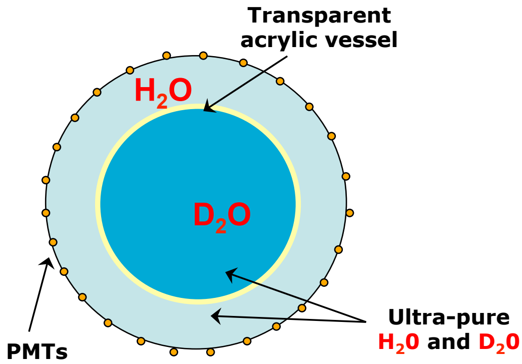
The Solar Neutrino "Problem"

Solar Neutrinos II: SNO

•Sudbury Neutrino Observatory located in a deep mine in Ontario, Canada



- 1000 ton heavy water (D_2O) Čerenkov detector
- D_2O inside a 12m diameter acrylic vessel
- Surrounded by 3000 tons of normal water
- Main experimental challenge is the need for very low background from radioactivity
- Ultra-pure H_2O and D_2O
- Surrounded by 9546 PMTs

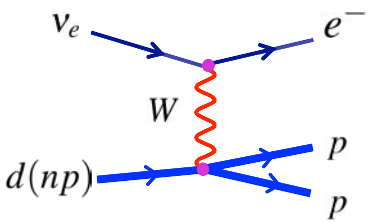


★ Detect Čerenkov light from three different reactions:

CHARGE CURRENT

- Detect Čerenkov light from electron
- Only sensitive to ν_e (thresholds)
- Gives a measure of ν_e flux

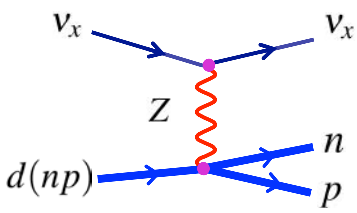
$$\boxed{\text{CC Rate} \propto \phi(\nu_e)}$$



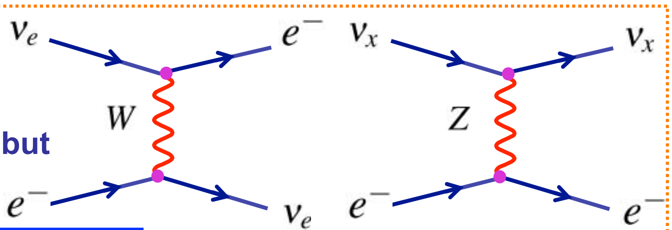
NEUTRAL CURRENT

- Neutron capture on a deuteron gives 6.25 MeV
- Detect Čerenkov light from electrons scattered by γ
- Measures total neutrino flux

$$\boxed{\text{NC Rate} \propto \phi(\nu_e) + \phi(\nu_\mu) + \phi(\nu_\tau)}$$



ELASTIC SCATTERING

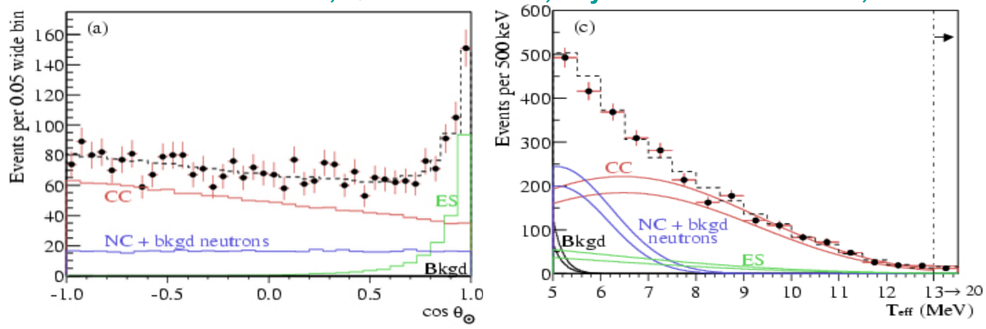


- Detect Čerenkov light from electron
- Sensitive to all neutrinos (NC part) – but larger cross section for ν_e

$$\boxed{\text{ES Rate} \propto \phi(\nu_e) + 0.154(\phi(\nu_\mu) + \phi(\nu_\tau))}$$

★ **Experimentally can determine rates for different interactions from:**
- **angle with respect to sun: electrons from ES point back to sun**
- **energy: NC events have lower energy – 6.25 MeV photon from neutron capture**
- **radius from centre of detector: gives a measure of background from neutrons**

SNO Collaboration, Q.R. Ahmad et al., Phys. Rev. Lett. 89:011301, 2002



★ **Using different distributions obtain a measure of numbers of events of each type:**

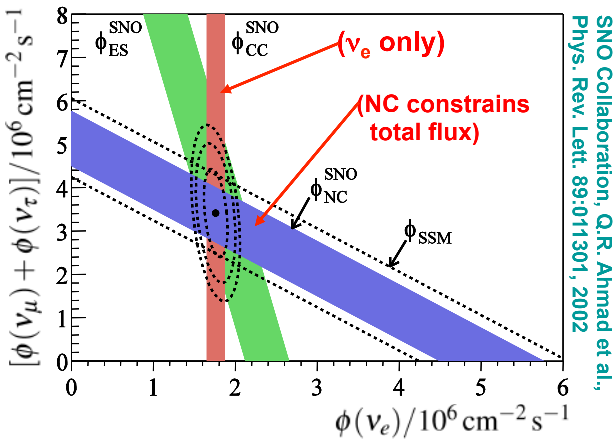
CC : 1968 ± 61	$\propto \phi(\nu_e)$
ES : 264 ± 26	$\propto \phi(\nu_e) + 0.154[\phi(\nu_\mu) + \phi(\nu_\tau)]$
NC : 576 ± 50	$\propto \phi(\nu_e) + \phi(\nu_\mu) + \phi(\nu_\tau)$

➡ **Measure of electron neutrino flux + total flux !**

★ **Using known cross sections can convert observed numbers of events into fluxes**

★ **The different processes impose different constraints**

★ **Where constraints meet gives separate measurements of ν_e and $\nu_\mu + \nu_\tau$ fluxes**



SNO Results:

$$\phi(\nu_e) = (1.8 \pm 0.1) \times 10^{-6}\,\mathrm{cm^{-2}\,s^{-1}}$$

$$\phi(\nu_\mu) + \phi(\nu_\tau) = (3.4 \pm 0.6) \times 10^{-6}\,\mathrm{cm^{-2}\,s^{-1}}$$

SSM Prediction:

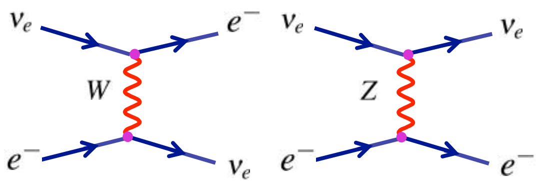
$$\phi(\nu_e) = 5.1 \times 10^{-6}\,\mathrm{cm^{-2}\,s^{-1}}$$

- **Clear evidence for a flux of ν_μ and/or ν_τ from the sun**
- **Total neutrino flux is consistent with expectation from SSM**
- **Clear evidence of $\nu_e \to \nu_\mu$ and/or $\nu_e \to \nu_\tau$ neutrino transitions**

Interpretation of Solar Neutrino Data

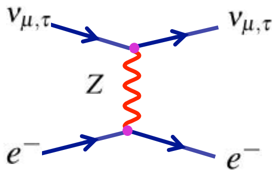
★ The interpretation of the solar neutrino data is complicated by **MATTER EFFECTS**

- The quantitative treatment is non-trivial and is not given here
- Basic idea is that as a neutrino leaves the sun it crosses a region of high electron density
- The coherent forward scattering process ($\nu_e \rightarrow \nu_e$) for an electron neutrino



CC + NC

is different to that for a muon or tau neutrino



NC

- Can enhance oscillations – "MSW effect"

★ A combined analysis of all solar neutrino data gives:

$$\Delta m^2_{\mathrm{solar}} \approx 8 \times 10^{-5}\,\mathrm{eV}^2,\ \sin^2 2\theta_{\mathrm{solar}} \approx 0.85$$

3) Reactor Experiments

- To explain reactor neutrino experiments we need the full three neutrino expression for the **electron neutrino survival probability (11)** which depends on U_{e1}, U_{e2}, U_{e3}
- Substituting these PMNS matrix elements in Equation (11):

$$\begin{aligned} P(\nu_e \rightarrow \nu_e) &\approx 1 - 4U_{e1}^2 U_{e2}^2 \sin^2\Delta_{21} - 4(1-U_{e3}^2)U_{e3}^2 \sin^2\Delta_{32} \\ &= 1 - 4(c_{12}c_{13})^2(s_{12}c_{13})^2 \sin^2\Delta_{21} - 4(1-s_{13}^2)s_{13}^2 \sin^2\Delta_{32} \\ &= 1 - c_{13}^4(2s_{12}c_{12})^2 \sin^2\Delta_{21} - (2c_{13}s_{13})^2 \sin^2\Delta_{32} \\ &= 1 - \cos^4\theta_{13} \sin^2 2\theta_{12} \sin^2\Delta_{21} - \sin^2 2\theta_{13} \sin^2\Delta_{32} \end{aligned}$$

- Contributions with short wavelength (atmospheric) and long wavelength (solar)
- For a 1 MeV neutrino

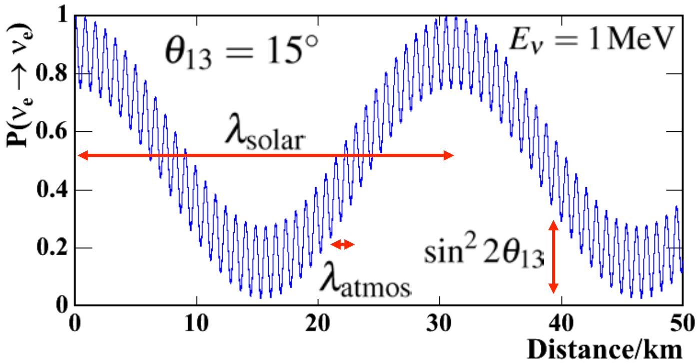
$$\lambda_{\mathrm{osc}}(\mathrm{km}) = 2.47 \frac{E(\mathrm{GeV})}{\Delta m^2(\mathrm{eV}^2)}$$

$$\Rightarrow \quad \lambda_{21} = 30.0\,\mathrm{km}$$

$$\lambda_{32} = 0.8\,\mathrm{km}$$

- Amplitude of short wavelength oscillations given by

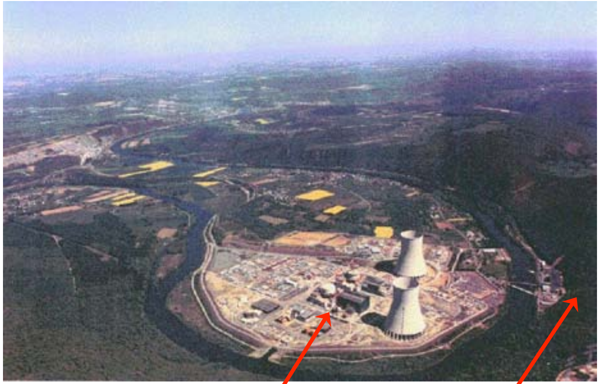
$$\sin^2 2\theta_{13}$$



Reactor Experiments I : CHOOZ

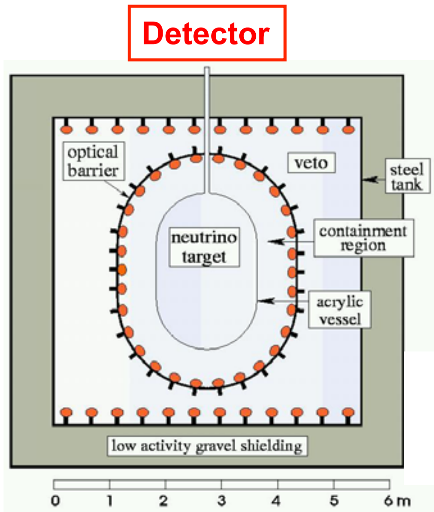
- Two nuclear reactors, each producing 4.2 GW
- Place detector 1km from reactor cores
- Reactors produce intense flux of $\overline{\nu}_e$

France



reactors

Detector
150m underground



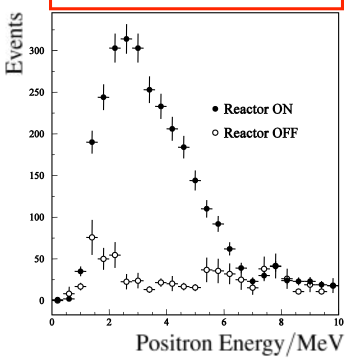
- Anti-neutrinos interact via inverse beta decay $\overline{\nu}_e + p \rightarrow e^+ + n$
- Detector is liquid scintillator loaded with Gadolinium (large n capture cross section)
- Detect photons from positron annihilation and a delayed signal from photons from neutron capture on Gadolinium

$$e^+ + e^- \rightarrow \gamma + \gamma$$
$$n + \mathrm{Gd} \rightarrow \mathrm{Gd}^* \rightarrow \mathrm{Gd} + \gamma + \gamma + \ldots$$

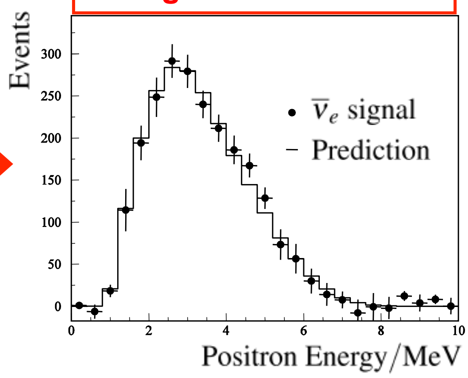
- At 1km and energies > 1 MeV, only the short wavelength component matters

$$P(\nu_e \rightarrow \nu_e) = 1 - \cos^4\theta_{13}\sin^2 2\theta_{12}\sin^2\Delta_{21} - \sin^2 2\theta_{13}\sin^2\Delta_{32} \approx 1 - \sin^2 2\theta_{13}\sin^2\Delta_{32}$$

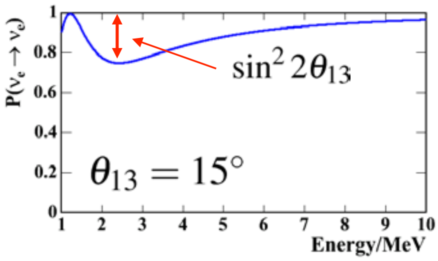
CHOOZ Raw Data



Background subtracted



Compare to effect of oscillations



★ Data agree with unoscillated prediction both in terms of rate and energy spectrum

$$N_{data}/N_{expect} = 1.01 \pm 0.04$$

CHOOZ Collaboration,
M.Apollonio et al.,
Phys. Lett. B420, 397-404, 1998

★ Hence $\sin^2 2\theta_{13}$ must be small !

$$\Rightarrow \sin^2 2\theta_{13} < 0.12 - 0.2$$

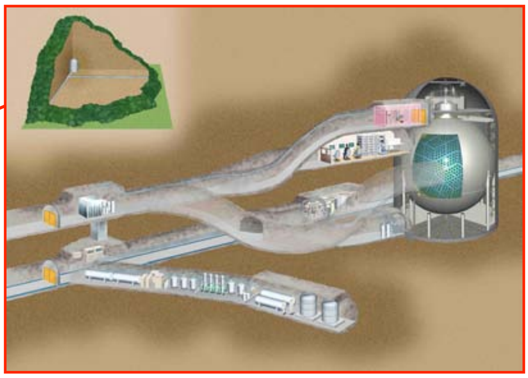
Exact limit depends on $|\Delta m^2_{32}|$

★ From atmospheric neutrinos (see appendix) can exclude $\theta_{13} \sim \frac{\pi}{2}$

- Hence the CHOOZ limit: $\sin^2 2\theta_{13} < 0.2$ can be interpreted as $\sin^2\theta_{13} < 0.05$

Reactor Experiments II : KamLAND

• Detector located in same mine as Super Kamiokande



18m

- 70 GW from nuclear power (7% of World total) from reactors within **130-240 km**
- Liquid scintillator detector, 1789 PMTs
- Detection via inverse beta decay: $\nu_e + p \to e^+ + n$

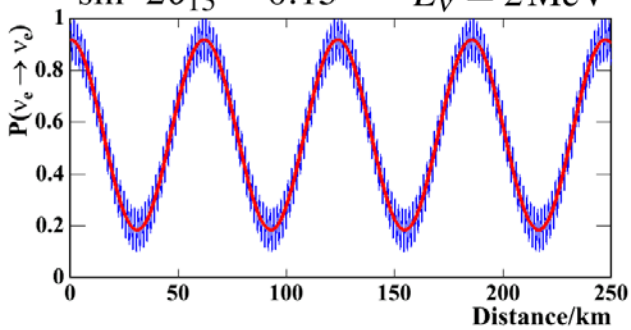
  Followed by $e^+ + e^- \to \gamma + \gamma$ prompt

  $n + p \to d + \gamma(2.2\,\mathrm{MeV})$ delayed

- For MeV neutrinos at a distance of 130-240 km oscillations due to Δm^2_{32} are very rapid
- Experimentally, only see average effect

$$\langle \sin^2 \Delta_{32} \rangle = 0.5$$

$\sin^2 2\theta_{13} = 0.15 \qquad E_\nu = 2\,\mathrm{MeV}$



★ Here:

$$\begin{aligned} P(\nu_e \to \nu_e) &= 1 - \cos^4\theta_{13}\sin^2 2\theta_{12}\sin^2\Delta_{21} - \sin^2 2\theta_{13}\sin^2\Delta_{32} \\ &\approx 1 - \cos^4\theta_{13}\sin^2 2\theta_{12}\sin^2\Delta_{21} - \frac{1}{2}\sin^2 2\theta_{13} \\ &= \cos^4\theta_{13} + \sin^4\theta_{13} - \cos^4\theta_{13}\sin^2 2\theta_{12}\sin^2\Delta_{21} \\ &\approx \cos^4\theta_{13}(1 - \sin^2 2\theta_{12}\sin^2\Delta_{21}) \end{aligned}$$

Averaging over rapid oscillations

neglect $\sin^4\theta_{13}$

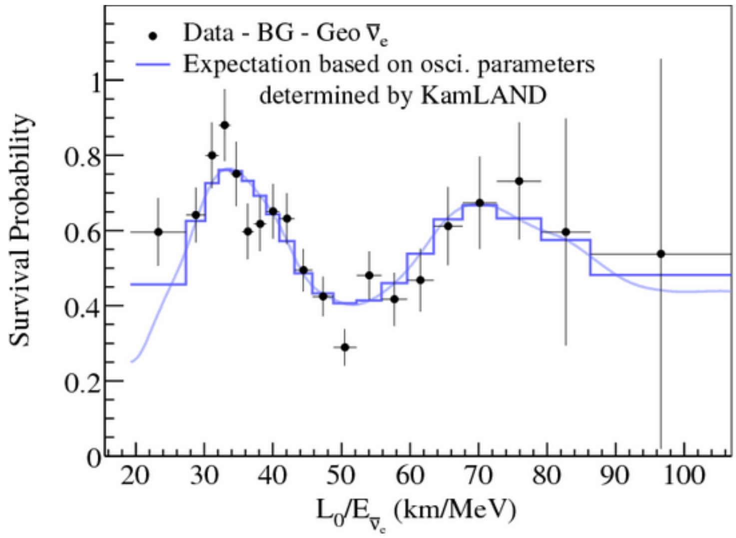
- Obtain two-flavour oscillation formula multiplied by $\cos^4\theta_{13}$
- From CHOOZ $\cos^4\theta_{13} > 0.9$

(Try Question 21)

KamLAND RESULTS:

Observe: **1609** events
Expect: **2179±89** events (if no oscillations)

KamLAND Collaboration, Phys. Rev. Lett., 221803, 2008



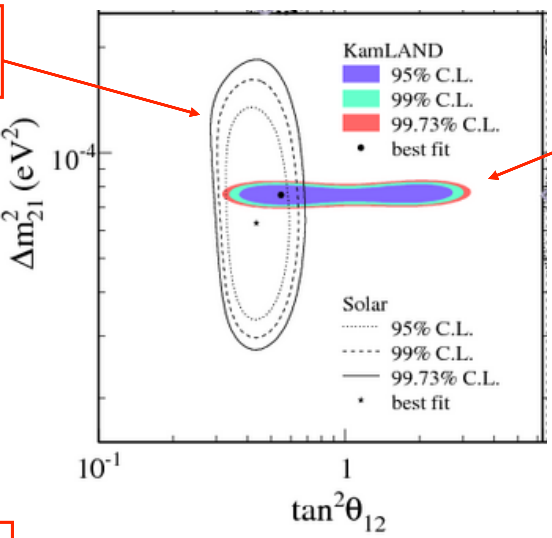
- ★ Clear evidence of electron anti-neutrino oscillations consistent with the results from solar neutrinos
- ★ Oscillatory structure clearly visible
- ★ Compare data with expectations for different osc. parameters and perform χ^2 fit to extract measurment

Combined Solar Neutrino and KamLAND Results

- ★ KamLAND data provides strong constraints on $|\Delta m^2_{21}|$
- ★ Solar neutrino data (especially SNO) provides a strong constraint on θ_{12}

Solar Neutrinos

KamLAND



Combined

$$|\Delta m^2_{21}| = (7.59 \pm 0.21)\times 10^{-5}\,\mathrm{eV}^2 \qquad \tan^2\theta_{12} = 0.47^{+0.06}_{-0.05}$$

STOP Press

★ In past few months, increasing evidence for non-zero value of non-zero θ_{13}

- T2K: $\nu_\mu \to \nu_e$ appearance (**2.5** σ)
- MINOS: $\nu_\mu \to \nu_e$ appearance (**2** σ) [Cambridge PhD work]
- Double-CHOOZ: $\overline{\nu}_e$ disappearance (**2** σ)

$$\boxed{\sin^2 2\theta_{13} \approx 0.04 - 0.08\,?}$$

Summary of Current Knowledge

SOLAR Neutrinos/KamLAND

KamLAND + Solar: $|\Delta m^2_{21}| \approx (7.6 \pm 0.2) \times 10^{-5}\,\mathrm{eV}^2$

SNO + KamLAND + Solar: $\tan^2\theta_{12} \approx 0.47 \pm 0.05$

➡ $\boxed{\sin\theta_{12} \approx 0.56; \quad \cos\theta_{12} \approx 0.82}$

Atmospheric Neutrinos/Long Baseline experiments

MINOS: $|\Delta m^2_{32}| \approx (2.4 \pm 0.1) \times 10^{-3}\,\mathrm{eV}^2$

Super Kamiokande: $\sin^2 2\theta_{23} > 0.92$

$$\boxed{\cos\theta_{23} \approx \sin\theta_{23} \approx \frac{1}{\sqrt{2}}}$$

CHOOZ + (atmospheric)

$$\boxed{\sin^2 2\theta_{13} < 0.15}$$

2011 hints

$$\boxed{\sin^2 2\theta_{13} \approx 0.04 - 0.08\,?}$$

★ Currently no knowledge about CP violating phase δ

- Regardless of uncertainty in θ_{13}

$$\begin{pmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{pmatrix} \approx \begin{pmatrix} c_{12} & s_{12} & ? \\ -s_{12}c_{23} & c_{12}c_{23} & s_{23} \\ s_{12}s_{23} & -c_{12}s_{23} & c_{23} \end{pmatrix}$$

- For the approximate values of the mixing angles on the previous page obtain:

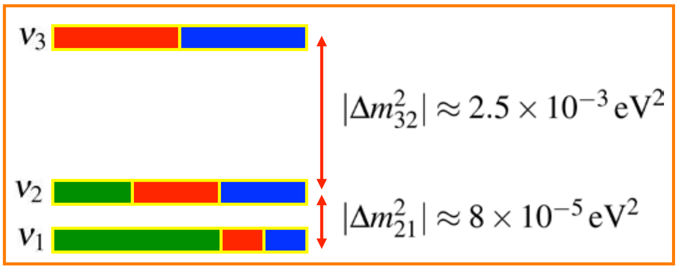
$$\begin{pmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{pmatrix} \approx \begin{pmatrix} 0.85 & 0.53 & 0.1e^{i\delta}? \\ -0.37 & 0.60 & 0.71 \\ 0.37 & -0.60 & 0.71 \end{pmatrix}$$

★ Have approximate expressions for mass eigenstates in terms of weak eigenstates:

$$|\nu_3\rangle \approx \tfrac{1}{\sqrt{2}}(|\nu_\mu\rangle + |\nu_\tau\rangle)$$

$$|\nu_2\rangle \approx 0.53|\nu_e\rangle + 0.60(|\nu_\mu\rangle - |\nu_\tau\rangle)$$

$$|\nu_1\rangle \approx 0.85|\nu_e\rangle - 0.37(|\nu_\mu\rangle - |\nu_\tau\rangle)$$



Final Words: Neutrino Masses

- Neutrino oscillations require non-zero neutrino masses
- But only determine mass-squared differences – not the masses themselves
- No direct measure of neutrino mass – only mass limits:

$$m_\nu(e) < 2\,\mathrm{eV}; \quad m_\nu(\mu) < 0.17\,\mathrm{MeV}; \quad m_\nu(\tau) < 18.2\,\mathrm{MeV}$$

Note the e, μ, τ refer to charged lepton flavour in the experiment, e.g. $m_\nu(e) < 2\,\mathrm{eV}$ refers to the limit from tritium beta-decay

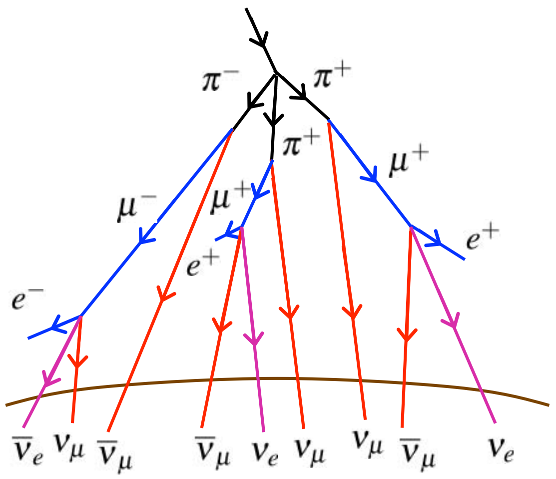
- Also from cosmological evolution infer that the sum

$$\boxed{\sum_i m_{\nu_i} < \text{few eV}}$$

★ 10 years ago – assumed massless neutrinos + hints that neutrinos might oscillate !
★ Now, know a great deal about massive neutrinos
★ But many unknowns: θ_{13}, δ, mass hierarchy, absolute values of neutrino masses
★ Measurements of these SM parameters is the focus of the next generation of expts.

Examinable Appendix: Atmospheric Neutrinos

- **High energy cosmic rays** (up to 10^{20} eV) interact in the upper part of the Earth's atmosphere
- The cosmic rays (~86% protons, 11% He Nuclei, ~1% heavier nuclei, 2% electrons) mostly interact hadronically giving showers of mesons (mainly pions)



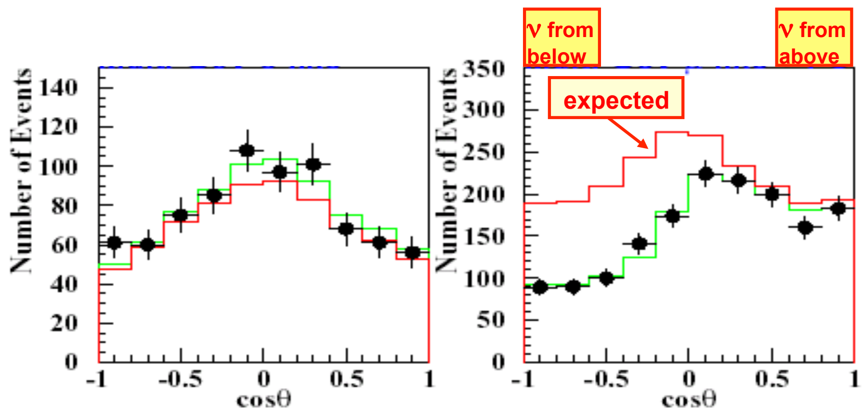
- Neutrinos produced by:

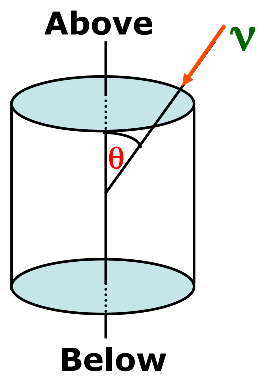
$$\pi^+ \to \mu^+ \nu_\mu \qquad \pi^- \to \mu^- \overline{\nu}_\mu$$
$$\hookrightarrow e^+ \nu_e \overline{\nu}_\mu \qquad \hookrightarrow e^- \overline{\nu}_e \nu_\mu$$

- Flux $\sim 1\,\mathrm{cm}^{-2}\mathrm{sr}^{-1}\mathrm{s}^{-1}$
- Typical energy : $\mathrm{E}_\nu \sim 1\,\mathrm{GeV}$
- Expect $\dfrac{N(\nu_\mu + \overline{\nu}_\mu)}{N(\nu_e + \overline{\nu}_e)} \approx 2$
- Observe a lower ratio with deficit of $\nu_\mu/\overline{\nu}_\mu$ coming from below the horizon, i.e. large distance from production point on other side of the Earth

Super Kamiokande Atmospheric Results

- Typical energy: $\mathrm{E}_\nu \sim 1\,\mathrm{GeV}$ (much greater than solar neutrinos – no confusion)
- Identify ν_e and ν_μ interactions from nature of Čerenkov rings
- Measure rate as a function of angle with respect to local vertical
- Neutrinos coming from above travel **~20 km**
- Neutrinos coming from below (i.e. other side of the Earth) travel **~12800 km**





★ Prediction for ν_e rate agrees with data

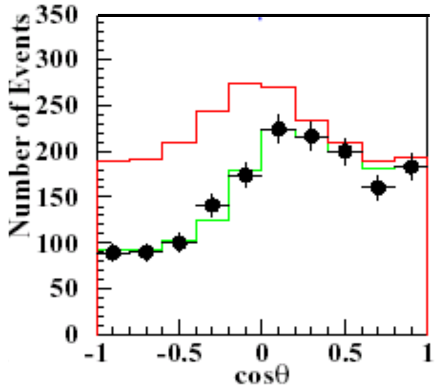
★ Strong evidence for disappearance of ν_μ for large distances

★ Consistent with $\nu_\mu \to \nu_\tau$ oscillations

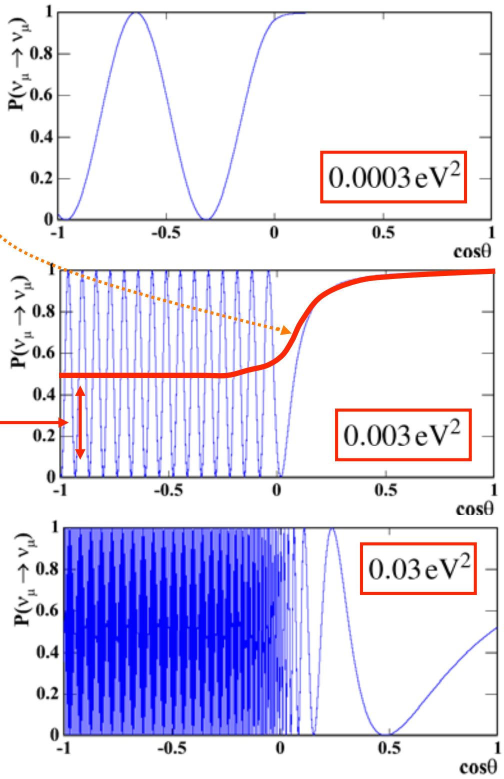
★ Don't detect the oscillated ν_τ as typically below interaction threshold of **3.5 GeV**

Interpretation of Atmospheric Neutrino Data

- Measure muon direction and energy not neutrino direction/energy
- Don't have E/θ resolution to see oscillations
- Oscillations "smeared" out in data
- Compare data to predictions for $|\Delta m^2|$



$$1 - \frac{1}{2}\sin^2 2\theta$$



★ Data consistent with:

$$|\Delta m^2_{\mathrm{atmos}}| \approx 0.0025\,\mathrm{eV}^2$$
$$\sin^2 2\theta_{\mathrm{atmos}} \approx 1$$

3-Flavour Treatment of Atmospheric Neutrinos

non-examinable

★ The energies of the detected atmospheric neutrinos are of order 1 GeV

★ The wavelength of oscillations associated with $|\Delta m^2_{21}| = 8\times 10^{-5}\,\mathrm{eV}^2$ is

$$\lambda_{21} = 31000\,\mathrm{km}$$

$$\lambda_{\mathrm{osc}}(\mathrm{km}) = 2.47\frac{E(\mathrm{GeV})}{\Delta m^2(\mathrm{eV}^2)}$$

- If we neglect the corresponding term in the expression for $P(\nu_\mu \to \nu_\tau)$ - equation (16)

$$\begin{aligned} P(\nu_\mu \to \nu_\tau) &\approx -4U_{\mu 1}U_{\tau 1}U_{\mu 2}U_{\tau 2}\sin^2\Delta_{21} + 4U^2_{\mu 3}U^2_{\tau 3}\sin^2\Delta_{32} \\ &\approx 4U^2_{\mu 3}U^2_{\tau 3}\sin^2\Delta_{32} \\ &= 4\sin^2\theta_{23}\cos^2\theta_{23}\cos^4\theta_{13}\sin^2\Delta_{32} \\ &= \cos^4\theta_{13}\sin^2 2\theta_{23}\sin^2\Delta_{32} \end{aligned}$$

- The Super-Kamiokande data are consistent with $\nu_\mu \to \nu_\tau$ which excludes the possibility of $\cos^4\theta_{13}$ being small
- Hence the CHOOZ limit: $\sin^2 2\theta_{13} < 0.2$ can be interpreted as $\sin^2\theta_{13} < 0.05$

NOTE: the three flavour treatment of atmospheric neutrinos is discussed below. The oscillation parameters in nature conspire in such a manner that the two flavour treatment provides a good approximation of the observable effects of atmospheric neutrino oscillations

3-Flavour Treatment of Atmospheric Neutrinos

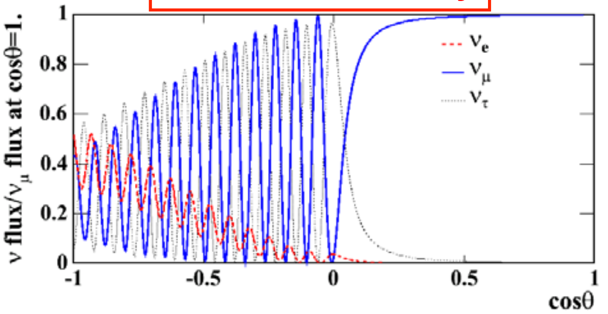
non-examinable

- Previously stated that the long-wavelength oscillations due to Δm^2_{21} have little effect on atmospheric neutrino oscillations because for a the wavelength for a 1 GeV neutrino is approx 30000 km.
- However, maximum oscillation probability occurs at $\lambda/2$
- This is not small compared to diameter of Earth and cannot be neglected
- As an example, take the oscillation parameters to be

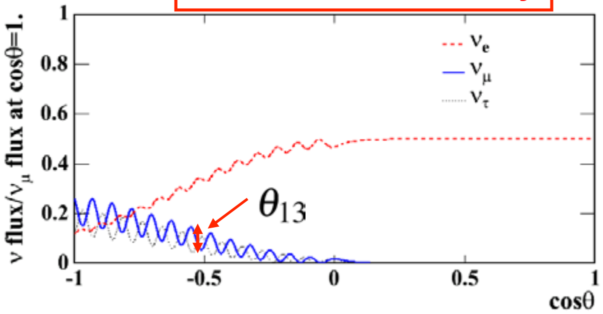
$$\theta_{12} = 32^\circ;\ \theta_{23} = 45^\circ;\ \theta_{13} = 7.5^\circ$$

- Predict neutrino flux as function of $\cos\theta$
- Consider what happens to muon and electron neutrinos separately

muon neutrinos only



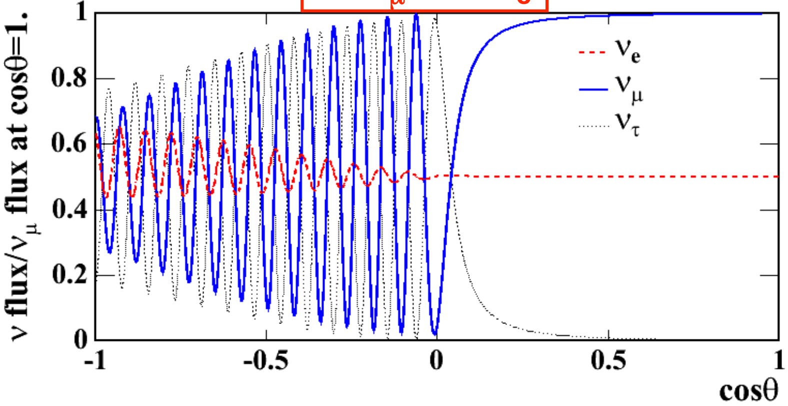
electron neutrinos only



- Δm^2_{21} has a big effect at $\cos\theta \sim -1$

- From previous page it is clear that the two neutrino treatment of oscillations of atmospheric muon neutrinos is a very poor approximation
- However, in atmosphere produce two muon neutrinos for every electron neutrino
- Need to consider the combined effect of oscillations on a mixed “beam” with both ν_μ and ν_e

2/3 ν_μ + 1/3 ν_e



- At large distances the average muon neutrino flux is still approximately half the initial flux, but only because of the oscillations of the original electron neutrinos and the fact that $\sin^2 2\theta_{23} \sim 1$
- Because the atmospheric neutrino experiments do not resolve fine structure, the observable effects of oscillations approximated by two flavour formula